

# Hierarchical star formation: stars and stellar clusters in the Gould Belt

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## ABSTRACT

We perform a study of the spatial and kinematical distribution of young open clusters in the solar neighbourhood, discerning between bound clusters and transient stellar condensations within our sample. Then, we discriminate between Gould Belt (GB) and local Galactic disc (LGD) members, using our previous estimate of the structural parameters of both systems obtained from a sample of O-B6 *Hipparcos* stars. Single membership probabilities of the clusters are also calculated in the separation process. Using this classified sample, we analyse the spatial structure and the kinematic behaviour of the cluster system in the GB. The two star formation regions that dominate and give the GB its characteristic-inclined shape show a striking difference in their content of star clusters: while Ori OB1 is richly populated by open clusters, not a single one can be found within the boundaries of Sco OB2. This is mirrored in the velocity space, translating again into an abundance of clusters in the region of the kinematic space populated by the members of Ori OB1, and a marginal number of them associated with Sco OB2. We interpret all these differences by characterizing the Orion region as a cluster complex typically surrounded by a stellar halo, and the Sco-Cen region as an OB association in the outskirts of the complex. In the light of these results, we study the nature of the GB with respect to the optical segment of the Orion Arm, and we propose that the different content of star clusters, the different heights over the Galactic plane and the different residual velocities of Ori OB1 and Sco OB2 can be explained in terms of their relative position to the density maximum of the Local Arm in the solar neighbourhood. Although morphologically intriguing, the GB appears to be the result of our local and biased view of a larger star cluster complex in the Local Arm, that could be explained by the internal dynamics of the Galactic disc.

**Key words:** stars: early-type – stars: formation – open clusters and associations: general – solar neighbourhood.

## 1 INTRODUCTION

The Gould Belt (GB) was first discovered by John Herschel (1847) and Benjamin Gould (1879) as a system of bright stars inclined with respect to the plane of the Milky Way. For more than a century, many studies have been devoted to describing its structure and kinematical behaviour, as well as to proposing a reliable global scenario that would account for its origin (for an extensive review on the subject, see Pöppel 1997, 2001; Grenier 2004). Today, it is considered that, in the scope of the most recent theories of hierarchical star formation (Efremov 1978, 1995; Elmegreen et al. 2000; Elmegreen 2006), the GB is likely to be our closest giant star-forming complex (Comérón 2001).

The stellar component of this complex takes the shape of a planar distribution of bright and young OB stars inclined with respect to the Galactic plane (Lesh 1968; Stothers & Frogel 1974; Westin 1985). Most of the young OB associations in the solar neighbourhood are known to be part of the GB (Blaauw 1965; de Zeeuw et al. 1999; Elias, Cabrera-Caño & Alfaro 2006a, hereafter Paper I). Also, a system of young, low-mass stars, detected by cross-matching X-ray and optical *Hipparcos* (Perryman 1997) based catalogues, appears to be associated with the GB (Guillout et al. 1998).

As we should expect from a giant ongoing star-forming complex, the local interstellar medium is prominently associated with the GB. The work by van den Bergh (1966) on reflection nebulae, by Sandqvist (1977) on dark clouds and, more recently, by Gaustad & Van Buren (1993) on maps of infrared emission have found a spatial distribution of the dark clouds of interstellar dust compatible with the pattern shown by the stellar component of the GB. Tomita (1986, 1987) and the study of the CO molecule by Taylor, Dickman

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& Scoville (1987) seem to confirm this inclination for the local molecular clouds.

Neutral hydrogen in the solar neighbourhood has also been related to the GB after the discovery of Lindblad's 'feature A' (Lindblad 1967; Lindblad et al. 1973), interpreted as a ring of gas with an expansion movement (Olano 1982; Elmegreen 1982). After Lindblad's work, the giant molecular clouds were found to be related to the most prominent OB star associations (Sancisi et al. 1974; Kutner et al. 1977; de Geus 1992). A full map of the CO molecule over the sky later confirmed that most of these clouds within 1 kpc from the Sun follow the GB pattern (Dame et al. 1993).

Then, if the GB is a star formation complex composed of both young stars and associations and interstellar material, we would expect to find a population of young open clusters following the GB pattern. This is obvious if we consider the concept of young star cluster in its widest sense, i.e. representing the dense inner regions of the hierarchical structure of young star fields (Elmegreen 2006). None the less, in this study we want to distinguish between a young cluster understood as a mere stellar condensation and a gravitationally bound system (i.e. as a condensation that 'has had sufficient time and gravitational self-attraction to get mixed by stellar orbital motions', as put by Elmegreen 2006). Our work focuses on the analysis of the cluster system, and on how these objects distribute and behave in comparison with the GB defined by the massive stars.

Since the first systematic studies that led to a discrimination between the GB and the LGD (Stothers & Frogel 1974; Westin 1985), there has been a great improvement in the number of catalogued open clusters, as well as in the availability of their astrometric and physical data. Surprisingly, it is not possible to find in the scientific literature any work devoted to the study of the open clusters membership to the GB and their distribution within this system before 2006. Thus, for many years, it has been accepted that the GB did not contain a significant population of bound clusters. Only very recently, Piskunov et al. (2006), in their analysis of the Galactic open cluster population, discovered an open cluster complex (OCC) that they associate with the GB. Although they find this OCC as a density peak within the spatial distribution of clusters in the solar neighbourhood, they estimate OCC membership probabilities by kinematical methods, through the analysis of the tangential velocities.

In two previous papers (Paper I; Elias, Alfaro & Cabrera-Caño 2006b, hereafter Paper II), we have studied the spatial distribution and the kinematic properties of the OB stars and associations in the GB. Our line of work thus leads in a logical way to the study of the distribution of young open clusters in the GB and their kinematic properties. Our analysis will be centred in the comparison of the GB morphology as obtained from the distribution of massive stars and clusters. This will represent another step in the understanding of the nature of the GB, and will also contribute to the knowledge of how star formation mechanisms proceed to the formation of stellar clusters.

## 2 OPEN CLUSTER SYSTEM

### 2.1 Associations, stellar condensations and bound clusters

Prior to describing the selection criteria for our observational sample, we want to punctuate some ideas about the concepts of hierarchical star formation, association, loose groupings and clusters as gravitationally bound physical systems.

Stars are born from molecular gas clouds whose internal structure can be characterized by a fractal dimension value that apparently

ranges from 2.5 to 2.7 (e.g. Sánchez, Alfaro & Pérez 2005, 2007). Simulations of the collapse of gas clouds hint at a hierarchical structure of the stellar formation, with clusters present in several dense cores (Walsh, Bourke & Myers 2006). During the past two decades, a set of observations has been collected which indicates that young stellar groupings show hierarchical patterns that range from the larger scales of flocculent spiral arms and star complexes to the smaller scales of OB associations, OB subgroups, small loose groups, clusters and cluster subclumps (Efremov 1995). In other words, the newly born stars seem to follow the same fractal pattern as the gas clouds from which they were originated.

The largest scale of stellar grouping, the great star complexes, would be associated with the gas superclouds with masses of the order of  $10^7$  solar masses. According to some authors (e.g. Efremov & Elmegreen 1998; Comerón 2001), the GB, with its spatial scale of the order of 1 kpc, would be the star complex closest to the Sun, and thus would come from a single gas cloud with a mass of a few million solar masses.

What happens at smaller scales? We already know that most of the OB associations in the solar neighbourhood, with typical sizes of about 80 pc, are mainly distributed along the plane of the GB. They represent the observable spatial scale immediately below the star complex. But, does it make any sense to talk about a typical scale within a fractal structure? The answer should be sought in the observational bias that is introduced by the age limit of the sample: OB associations are detected and selected as concentrations of OB stars with a typical age of about 10 Myr. The existence of some general correlation between the duration of the star formation and the size of the region (Larson 1981; Elmegreen 2006) implies that, for a typical age of 10 Myr, the typical size of the region is 80 pc (Efremov & Elmegreen 1998; Elmegreen 2006).

Star clusters are formed in the cores of giant molecular clouds; they represent the stellar groupings associated with the inner and densest regions of the gas, and can be interpreted as the unavoidable result of star formation in hierarchically structured gas (Elmegreen 2006). However, only a few of these condensations will still be gravitationally bound after 10 Myr. It has been estimated that 90 per cent of the clusters lose a high fraction of their stars in the first 10 Myr of their lives (Fall, Chandar & Whitmore 2005). Thus, after 10 Myr of life, it is possible to find star-forming regions that maintain a large number of star clusters, while others like NGC 604 in M33 do not contain a single cluster (Maíz-Apellániz 2001). This seems to depend on variations of the mean density of the clouds; those where the average density is low will form stellar concentrations in the cores of the clouds, but they will not have enough binding energy to keep a bound cluster when the gas leaves. Thus, the clusters observed within a star-forming region at a certain moment could be representatives of two distinct physical systems: either star clusters, gravitationally bound and able to survive galactic tidal forces, or a mere stellar condensation with a mean lifetime of the order of 10 million years or less.

Thus, the analysis of the distribution of star clusters (both bound systems and transient condensations) could give us information about the history of star formation in the GB, as well as about the physical conditions of the gas from which they were born.

### 2.2 Cluster sample

We extract our sample from the Catalogue of Open Cluster Data (COCD) and its Extension 1, compiled by Kharchenko et al. (2005a,b). This catalogue has the advantage of homogeneity over other existing compilations, and since we are also

interested in working with space velocities, the fact that the COCD had catalogued proper motion values in the *Hipparcos* system and newly determined radial velocity data was decisive in our choice.

As our aim is to study the cluster distribution in the GB, we establish distance and age limits in our sample. The GB system should be well contained within a heliocentric radius of 1 kpc (Stothers & Frogel 1974; Westin 1985; Paper I); and since its age has been estimated between 20 and 90 Myr (Torra, Fernández & Figueras 2000), we only keep clusters younger than 100 Myr.

Thus, our first selection (that will be reduced after outlier elimination, as we explain in the following section) is composed of 93 open clusters, 83 of which have complete kinematical information. We calculate for every cluster its Cartesian Galactic coordinates ( $X, Y, Z$ ), where  $X$  is positive in the direction of the Galactic centre,  $Y$  in the direction of Galactic rotation and  $Z$  perpendicular to the Galactic plane so that they form a right-handed, orthogonal frame. We also calculate their respective space velocities, ( $U, V, W$ ), for those clusters with radial velocity data.

### 2.3 Gould Belt and Local Galactic Disc clusters

In Paper I, we developed a three-dimensional classification method that allowed us to separate the GB stars from the local Galactic disc (LGD) stars by purely spatial criteria. This method considered that the LGD and the GB could be described as a distribution along two intersecting planes; working with a sample of *Hipparcos* O-B6 stars we obtained an estimation of the parameters that defined those planes. In Paper II, we used those parameters to classify strictly by spatial criteria a new sample of OB stars with kinematic data. Similarly, we now use our classification method and the parameters found in Paper I to separate the GB open clusters from those of the LGD field. We want to stress that this separation is done only by means of their spatial position in the three-dimensional space.

As described in Paper I, this classification method leads also to the identification of outliers; i.e. those objects that are too far away from the mean planes and thus lie in regions of low density of probability in the sample space. We have found nine outliers in our cluster sample; after their elimination the remaining sample consists of 84 clusters, 76 of them with full space velocity data (listed in Table 1). Further information on the detection and meaning of outliers may be found in section 2 of Paper I and references therein.

Finally, the individual GB membership probability ( $P$ ) for each cluster is listed in Table 1. In total, 40 clusters have a probability  $P > 50$  per cent, and thus, following the Bayes minimum error rate decision rule, we consider them to be members of the GB.

### 2.4 Bound and unbound clusters

This sample may contain both bound and unbound clusters, as we have previously seen (Section 2.1). A simple criterion to select those condensations with a high probability of being bound clusters is an age cut-off, if we consider that all the objects older than 10 Myr can be called bound clusters, in the sense that they have survived to the high rate of infant mortality that happens during the first 10 million years in the life of a cluster. Once this critical threshold has been surpassed, the mean life of a cluster, although very dependent on the environmental conditions, is usually larger than 1 Gyr.

It is evident, though, that such a selection is just an approximation to the problem, and that to determine if a stellar system

represents a bound cluster, we must compare its density with the tidal density at its position in the Galaxy. However, we lack the complete information to perform this detailed analysis, and thus we must resort to empirical classification criteria. The COCD catalogue provides three variables for each cluster that may give us some additional information about what we understand as young stellar grouping and cluster; these variables are cluster radius, core radius and age.

In Fig. 1, we represent the cluster sample in the cluster radius–age space. It is interesting to note that the clusters older than 10 Myr in our sample seem to group around an elongated strip with a positive slope and that only a few objects present a cluster radius that deviates from this mean distribution. The most notorious case is that of NGC 2264, which has a radius close to 20 pc. This cluster seems to present a high degree of substructure, spatial as well as kinematical (Fűrész et al. 2006), that its central potential has not been able to erase. Thus, we eliminate this object from the sample, for its properties appear to differ in some structural aspects from that of ‘classic’ bound clusters.

We have also drawn in Fig. 1 an upper envelope of the main cluster distribution with ages older than 10 Myr and younger than 100 Myr. There seems to be a natural separation between the main distribution of clusters older than 10 Myr and those that show a radius larger than expected for their age. Based on this apparent separation, we have traced this upper envelope of the main cluster distribution. The exact location of the line has been drawn by consensus of the authors. It simply aims to represent a dividing line between the largest concentration of ‘probable’ bound clusters (represented as filled circles in the figure) and those that in spite of their age show a clearly distinct behaviour. The objects located under this envelope line can be considered, from a conservative point of view, as highly probable bound clusters.

Using the cluster radius and core radius data, we have estimated a pseudo-concentration parameter in the form  $\log(R_{cl}/R_{co})$ , that we represent in Fig. 2 versus the cluster radius. Circled squares indicate those clusters that, according to the previously exposed criteria, can be considered as probable bound clusters. We note that most of these objects are distributed, as in Fig. 1, along a straight line. However, two of these objects (Platais 6 and NGS 2546) show a clear separation from the general tendency.

Using the online tools at WEBDA (WEB Base Données Amas) data base (Mermilliod 1995), we have recalculated the parameters of these two clusters from their photometric data. This analysis indicates that the photometric diagrams of Platais 6 show a good fitting for a distance modulus of 7 and  $\log(t) \approx 6.5$ . Similarly, a visual inspection of the colour–magnitude diagram of NGC 2546, as catalogued by WEBDA, seems to indicate that it is a very young cluster [ $\log(t) \leq 7$ ] with a rich population of pre-main-sequence stars. This diagram is similar to that shown by NGC 2362 and other young clusters located in the third Galactic quadrant (Delgado et al. 2006; Delgado, Alfaro & Yun 2007), meaning that we are not facing an object almost 100 Myr old. Thus, we consider these two objects as clusters with ages inferior to 10 Myr, and consequently they do not belong – according to our criteria – to the group of clusters with a high probability of being gravitationally bound.

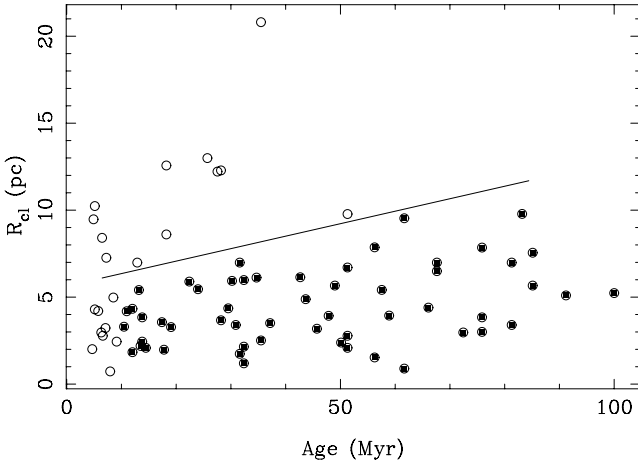
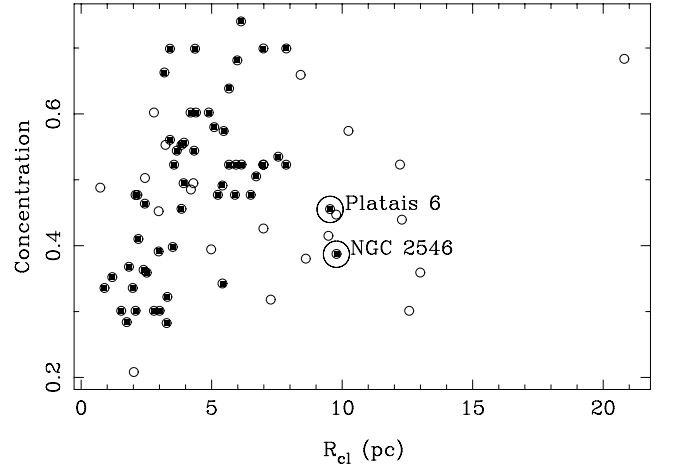
In Table 1, we present the classification of the sample that follows from this reasoning. The last column shows an indicator of the probability of being a bound cluster, according to the criteria previously discussed (1 stands for those objects that are ‘probable’ bound clusters, and 2 stands for those we consider as transient stellar condensations).

**Table 1.** Catalogue of stellar clusters used in this study. ( $X$ ,  $Y$ ,  $Z$ ) are the spatial coordinates of the clusters in the directions of Galactic centre, Galactic rotation and north Galactic pole. ( $U$ ,  $V$ ,  $W$ ) are their respective heliocentric space velocities, and ( $U'$ ,  $V'$ ,  $W'$ ) their residual velocities corrected from solar motion and Galactic differential rotation.  $R_{cl}$  and  $R_{co}$  stand for cluster radius and core radius, respectively.  $T$  is the cluster age.  $P$  is the GB membership probability calculated with our method exclusively by spatial criteria. *Class* describes whether a cluster is probably bound (1) or unbound (2).

COCD	Name	$X$ (pc)	$Y$ (pc)	$Z$ (pc)	$U$ (km s <sup>-1</sup> )	$V$ (km s <sup>-1</sup> )	$W$ (km s <sup>-1</sup> )	$U'$ (km s <sup>-1</sup> )	$V'$ (km s <sup>-1</sup> )	$W'$ (km s <sup>-1</sup> )	$R_{cl}$ (pc)	$R_{co}$ (pc)	$\log(T)$ (yr)	$P$ (per cent)	<i>Class</i>
2	Berkeley 59	-471	878	87	20.8	3.1	6.8	2.0	15.7	13.1	3.0	1.0	6.80	10	2
32	Stock 7	-491	496	1	10.5	9.7	-7.9	3.9	22.3	-1.6	2.2	0.9	7.13	7	1
35	Trumpler 2	-476	438	-45	22.1	-34.9	-11.3	17.3	-22.3	-5.0	5.7	1.7	7.93	54	1
41	Stock 23	-291	244	14	13.4	-9.9	-3.6	14.9	2.7	2.7	1.2	0.5	7.51	17	1
42	Melotte 20	-159	102	-21	-12.9	-26.4	-7.5	-6.8	-13.8	-1.2	20.8	4.3	7.55	50	2
46	IC 348	-354	125	-121	-16.9	-11.6	-8.3	-11.6	1.0	-2.0	0.9	0.4	7.79	94	1
68	Collinder 65	-301	-47	-58	-14.8	-11.3	-6.9	-4.0	1.3	-0.6	13.0	5.7	7.41	67	2
72	Collinder 69	-413	-113	-92	-28.3	-11.7	-7.9	-15.4	0.9	-1.6	4.2	1.4	6.76	84	2
73	NGC 1981	-334	-178	-130	-24.6	-11.9	-6.5	-9.6	0.7	-0.2	1.7	0.9	7.50	95	1
74	NGC 1976	-329	-183	-132	-23.2	-16.3	-7.1	-8.1	-3.7	-0.8	2.8	1.4	7.71	95	1
75	NGC 1977	-415	-225	-164	-18.7	-17.7	-5.4	-2.2	-5.1	0.9	1.8	0.8	7.08	97	1
76	NGC 1980	-451	-255	-184	-20.3	-13.6	-7.1	-2.9	-1.0	-0.8	2.0	1.2	6.67	98	2
77	Collinder 70	-338	-158	-117	-16.3	-9.2	-5.8	-1.9	3.4	0.5	10.2	2.7	6.71	95	2
80	Sigma Ori	-340	-172	-119	-25.1	-16.4	-3.5	-10.3	-3.8	2.8	2.8	0.7	6.82	95	2
91	Platais 6	-313	-148	-38	-21.2	-13.1	-12.5	-7.2	-0.5	-6.2	9.5	3.3	7.79	27	2
93	Collinder 89	-789	-120	53	-24.8	-11.8	-0.7	-11.7	0.8	5.6	7.0	2.1	7.50	0	1
95	NGC 2232	-265	-183	-42	-12.6	-9.0	-9.7	2.5	3.6	-3.4	3.4	0.7	7.49	38	1
107	NGC 2264	-607	-257	25	-14.4	-12.5	-11.8	3.1	0.1	-5.5	8.4	1.8	6.81	1	2
125	Alessi 21	-363	-344	0	-35.3	-19.1	-4.3	-15.0	-6.5	2.0	4.4	0.9	7.47	6	1
126	Collinder 132	-183	-362	-66	-24.5	-17.5	-10.4	-3.6	-4.9	-4.1	2.2	0.7	7.51	74	1
133	Collinder 135	-112	-292	-62	-17.9	-7.0	-13.9	0.7	5.6	-7.6	6.1	1.1	7.54	82	1
136	Collinder 140	-168	-361	-55	-21.3	-13.4	-14.3	-0.4	-0.8	-8.0	3.5	1.4	7.57	62	1
143	Bochum 4	-528	-694	12	-13.8	-9.7	-4.9	17.7	2.9	1.4	2.0	0.9	7.25	1	1
155	Haffner 13	-301	-646	-46	-52.5	-49.7	-14.4	-22.6	-37.1	-8.1	6.0	1.2	7.51	12	1
159	NGC 2451A	-56	-178	-25	-26.2	-12.8	-13.4	-11.2	-0.2	-7.1	5.4	2.5	7.76	53	1
162	NGC 2451B	-132	-406	-50	-20.4	-6.0	-15.3	1.9	6.6	-9.0	3.0	1.5	7.88	58	1
182	Vel OB2	-48	-404	-57	-22.6	-12.6	-3.0	-0.3	0.0	3.3	8.6	3.6	7.26	78	2
183	NGC 2547	-44	-450	-68	-18.9	-11.1	-13.8	4.8	1.5	-7.5	2.4	1.0	7.70	78	1
186	NGC 2546	-240	-886	-33	-37.4	-26.8	-9.3	0.2	-14.2	-3.0	9.8	4.0	7.92	7	2
190	vdBergh-Hagen 23	-120	-420	-8	-24.8	-10.9	-5.0	-2.1	1.7	1.3	2.4	0.8	7.14	25	1
202	IC 2391	1	-175	-21	-27.7	-15.6	-6.1	-12.8	-3.0	0.2	3.8	1.1	7.88	57	1
204	Mamajek 1	37	-90	-39	-13.5	-18.6	-10.6	-1.3	-6.0	-4.3	0.7	0.2	6.90	57	2
205	IC 2395	-42	-706	-44	-16.8	-23.6	-8.1	15.1	-11.0	-1.8	4.3	1.2	7.08	41	1
210	Trumpler 10	-53	-414	5	-25.8	-12.0	-10.6	-3.3	0.6	-4.3	5.5	1.5	7.38	35	1
213	vdBergh-Hagen 56	-65	-677	17	-23.3	20.3	0.3	7.7	32.9	6.6	3.6	1.1	7.24	26	1
216	Platais 8	20	-147	-20	-13.0	-18.8	-3.7	1.0	-6.2	2.6	7.9	1.6	7.75	62	1
255	vdBergh-Hagen 99	151	-509	-5	-28.7	-14.6	-16.2	-3.1	-2.0	-9.9	3.0	1.2	7.86	68	1
259	IC 2602	53	-150	-14	-8.2	-23.1	-0.4	5.9	-10.5	5.9	7.0	1.4	7.83	62	1
261	Alessi 5	123	-378	-14	-22.3	-20.5	-7.2	-0.9	-7.9	-0.9	2.1	0.7	7.71	66	1
357	vdBergh-Hagen 164	382	-392	-58	-13.4	-29.2	-19.1	8.5	-16.6	-12.8	3.8	1.3	7.14	13	1
366	NGC 6025	623	-444	-79	-12.5	-12.2	3.4	11.0	0.4	9.7	5.1	1.3	7.96	3	1
371	NGC 6087	758	-479	-85	-13.9	-4.7	-1.9	10.7	7.9	4.4	7.5	2.2	7.93	1	1
395	NGC 6322	961	-252	-53	-56.4	11.3	-6.8	-39.1	23.9	-0.5	2.1	1.0	7.16	0	1
402	NGC 6383	982	-74	1	3.6	-2.2	-10.5	15.3	10.4	-4.2	4.3	1.4	6.71	0	2
408	NGC 6405	486	-29	-6	-12.3	-12.1	-4.4	-2.1	0.5	1.9	3.4	0.9	7.91	6	1
412	IC 4665	290	171	103	-1.4	-14.2	-7.5	2.4	-1.6	-1.2	6.1	1.8	7.63	92	1
423	Collinder 359	542	310	140	10.4	-21.4	-13.6	9.8	-8.8	-7.3	12.3	4.5	7.45	99	2
425	NGC 6514	810	101	-4	-2.4	1.3	-10.8	3.7	13.9	-4.5	3.3	1.7	7.28	0	1
449	IC 4725	601	146	-48	2.2	-16.6	0.6	6.8	-4.0	6.9	6.5	2.2	7.83	2	1
456	Stephenson 1	142	331	99	-4.5	-19.4	-10.2	-5.8	-6.8	-3.9	5.7	1.3	7.69	83	1
479	Roslund 5	133	396	2	-5.8	-18.9	-6.9	-9.2	-6.3	-0.6	3.9	1.1	7.77	28	1
484	Collinder 419	153	723	36	23.2	-12.9	-6.7	9.3	-0.3	-0.4	3.2	0.9	6.85	47	2
500	IC 1396	-134	822	54	19.4	-11.9	-7.2	2.4	0.7	-0.9	9.5	3.6	6.69	56	2
501	NGC 7160	-191	766	89	19.0	-22.4	-1.0	3.8	-9.8	5.3	3.2	0.7	7.66	43	1
506	Pismis-Moreno 1	-258	858	83	6.4	-22.0	-4.3	-11.7	-9.4	2.0	2.5	1.1	7.55	34	1
510	Cep OB3	-253	651	36	12.0	-11.3	-4.8	0.4	1.3	1.5	12.2	3.7	7.44	31	2
1013	ASCC 13	-765	228	45	5.7	-5.3	-6.4	7.7	7.3	-0.1	9.8	3.5	7.71	1	2
1016	ASCC 16	-408	-156	-145	-0.6	-1.5	0.9	13.7	11.1	7.2	5.0	2.0	6.93	97	2
1018	ASCC 18	-439	-178	-159	-11.9	-6.0	-2.5	3.1	6.6	3.8	5.4	1.7	7.12	98	1

**Table 1** – *continued*

COCD	Name	X (pc)	Y (pc)	Z (pc)	U (km s <sup>-1</sup> )	V (km s <sup>-1</sup> )	W (km s <sup>-1</sup> )	U' (km s <sup>-1</sup> )	V' (km s <sup>-1</sup> )	W' (km s <sup>-1</sup> )	R <sub>cl</sub> (pc)	R <sub>co</sub> (pc)	log(T) (yr)	P (per cent)	Class
1019	ASCC 19	-299	-139	-117	-16.0	-9.0	-6.1	-2.3	3.6	0.2	4.9	1.2	7.64	94	1
1020	ASCC 20	-399	-158	-136	-16.1	-5.3	-5.0	-1.7	7.3	1.3	5.9	2.0	7.35	97	1
1021	ASCC 21	-451	-163	-142	-17.4	-8.3	-5.3	-2.9	4.3	1.0	7.0	2.6	7.11	98	2
1024	ASCC 24	-318	-236	-57	-10.2	-10.3	-15.1	6.6	2.3	-8.8	2.4	0.8	6.96	47	2
1033	ASCC 33	-434	-661	-121	-23.6	-1.0	-8.8	6.8	11.6	-2.5	12.6	6.3	7.26	81	2
1047	ASCC 47	-183	-881	5	-41.9	-2.6	-17.6	-4.4	10.0	-11.3	7.9	2.4	7.88	10	1
1050	ASCC 50	-111	-843	22	-30.4	-13.5	-9.6	5.9	-0.9	-3.3	5.9	1.8	7.48	17	1
1058	ASCC 58	122	-587	14	-33.0	-15.3	-13.0	-4.9	-2.7	-6.7	4.2	1.1	7.04	60	1
1069	ASCC 69	482	-867	-126	-29.8	-17.5	-9.4	7.2	-4.9	-3.1	7.0	2.1	7.91	11	1
1076	ASCC 76	376	-465	-44	-19.7	2.7	-5.4	4.5	15.3	0.9	3.7	1.1	7.45	15	1
1079	ASCC 79	612	-513	-40	-8.1	-15.3	-7.9	17.7	-2.7	-1.6	7.3	3.5	6.86	3	2
1084	ASCC 84	721	-532	-85	-16.2	-11.7	-10.7	10.2	0.9	-4.4	3.9	1.3	7.68	1	1
1104	ASCC 104	460	654	-22	9.3	-13.7	-8.2	-2.3	-1.1	-1.9	6.7	2.1	7.71	2	1
1105	ASCC 105	229	444	19	-2.5	-19.5	-7.0	-7.4	-6.9	-0.7	5.2	1.8	8.00	31	1
1114	ASCC 114	-68	546	10	15.5	-20.2	-0.8	7.4	-7.6	5.5	1.5	0.8	7.75	62	1
1118	ASCC 118	-224	869	66	27.4	-35.6	-11.6	8.9	-23.0	-5.3	3.3	1.6	7.02	41	1
1127	ASCC 127	-132	323	25	-5.5	-10.5	-8.4	-6.6	2.1	-2.1	4.4	1.1	7.82	42	1

**Figure 1.** Cluster radius versus age. The filled diamonds represent *our probable* bound clusters in the age range from 10 to 100 Myr.**Figure 2.** Pseudo-concentration parameter,  $\log(R_{cl}/R_{co})$  versus cluster radius. Filled diamonds represent *our probable* bound clusters.

### 3 ANALYSIS AND DISCUSSION

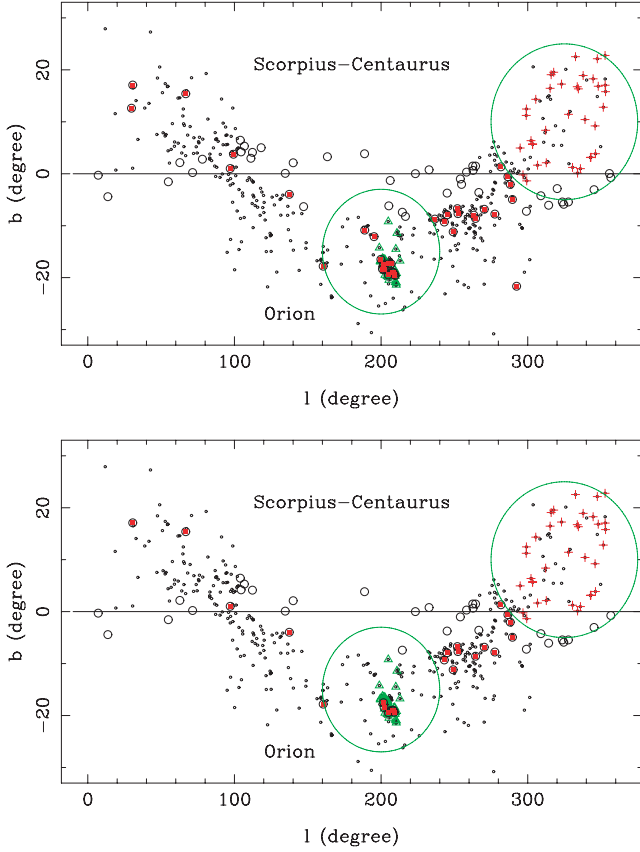
#### 3.1 Spatial distribution

A two-dimensional projection of the spatial distribution of open clusters in the sky is shown in Fig. 3 (top panel). There we see, as we commented in Section 1, that GB clusters (filled circles) are mostly concentrated towards southern Galactic latitudes, and that only three of them ( $\sharp 395 = \text{IC 4665}$ ,  $\sharp 423 = \text{Collinder 359}$  and  $\sharp 456 = \text{Stephenson 1}$ ) clearly rise above the Galactic plane ( $b > 10^\circ$ ). According to our criteria,  $\sharp 395$  and  $\sharp 456$  are probable bound clusters, while  $\sharp 423$  could rather be a transitory stellar condensation. We must note that these three objects are located in the first Galactic quadrant and that, apparently, they are not related to any OB association.

We also represent in this figure the OB stars used in our analysis of the spatial structure of the GB (Paper I). The associations Sco OB2, Ori OB1, Per OB2 and Lac OB1 are classically thought to be

components of the GB (e.g. Olano 1982), and studies that followed the *Hipparcos* mission suggest that Vel OB2, Tr 10 and Collinder 121 also belong to the GB (de Zeeuw et al. 1999), although their position close to the line of nodes where the GB intersects the Galactic disc adds quite some uncertainty to this membership assignment. In any case, it is evident that the Sco OB2 and Ori OB1 associations, respectively, mark the south and north Galactic extremes of the inclined plane which best describes the stellar system known as the GB, and that in a certain sense the geometrical characterization of the GB is defined by the relative position of these two associations.

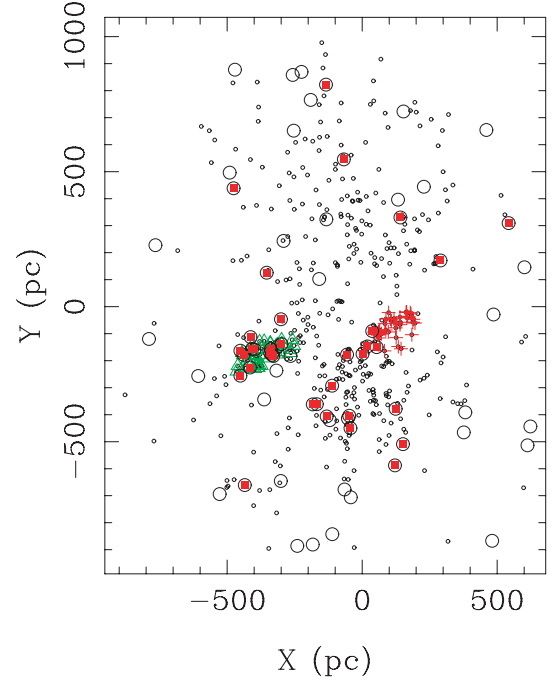
Thus, we have marked with a red cross the stars belonging to the Sco OB2 association, following the coordinates as given by de Zeeuw et al. (1999), that situate the complex in the range  $l = 290^\circ\text{--}360^\circ$ ,  $b = -10^\circ\text{--}30^\circ$  and  $D = 100\text{--}220$  pc, and that we have tagged as Scorpius-Centaurus in the figure. In the same manner, we have selected the stars belonging to Ori OB1 as those in the range  $l = 197^\circ\text{--}215^\circ$ ,  $b = -12^\circ\text{--}26^\circ$  and  $D = 300\text{--}550$  pc (de Zeeuw



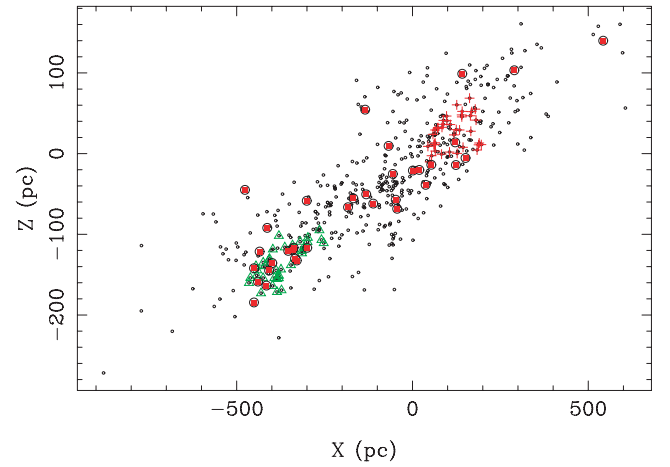
**Figure 3.** Distribution of open clusters on the sky (Galactic coordinates). Open and filled diamonds represent clusters classified as LGD and GB members, respectively. Small circles stand for the O-B6 GB stars from Paper I. Red crosses and green triangles mark, respectively, the stars belonging to Sco OB2 and Ori OB1, according to de Zeeuw et al. (1999). Top panel shows all the clusters in our sample, whereas bottom panel displays only the probable bound clusters.

et al. 1999). These stars appear in Fig. 3 as green triangles, and the region is tagged as Orion. Estimations of the age range of these associations indicate that Sco OB2 members are between 5 and 20 Myr old (Sartori, Lépine & Dias 2003), while the typical age of Ori OB1 members is between some  $10^5$  yr and 11 Myr (Briceño et al. 2005; Hernández et al. 2006).

Hence, we are facing two associations with sizes and ages relatively similar, and according to de Zeeuw et al. (1999), with a number of probable Sco OB2 members larger than that of Ori OB1. With these data, it sounds reasonable enough to think that the star cluster population related to Sco OB2 should be larger than that related to Ori OB1. But what we observe in Fig. 3 (top panel) is exactly the opposite. There is not a single cluster within the frontiers of Sco OB2, while we have detected 11 objects associated with Ori OB1 (#73, #74, #75, #76, #77, #80, #1016, #1018, #1019, #1020, #1021), six of them being ‘probable’ bound clusters (#73=NGC 1981, #74=NGC 1976, #75 = NGC 1977, #1018, #1019, #1020). This is better observed in Fig. 3 (bottom panel), where we have represented only those clusters catalogued as ‘probable’ bound ones, tagged with number 1 in the last column of Table 1. This figure shows how almost the totality of the ‘probable’ bound clusters associated with the GB are located in the Orion and the Puppis-Vela regions. Moreover, almost all of the other objects show no relationship with other OB associations in the GB.

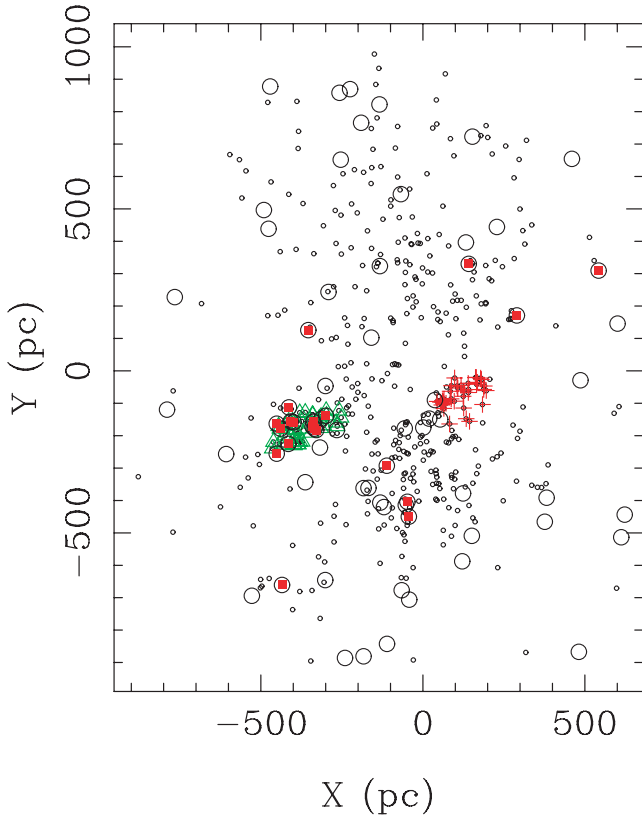


**Figure 4.** Spatial distribution of the sample clusters in the XY plane. Symbols are as in Fig. 3.



**Figure 5.** Spatial distribution of GB clusters (filled circles) in the XZ plane. Other symbols are as in Fig. 3.

The same phenomenology can be observed in Figs 4 and 5, where we have represented the GB cluster distribution in the spatial projections XY and XZ. The LGD clusters tend to distribute uniformly across the Galactic plane in Fig. 4, but that is not the case for GB open clusters. We observe how the GB’s distribution is quite clumpy; many of the clusters are located in the region of Vela, and we specially note that the region of Orion (in the third quadrant of the XY plane, and around the most negative values of Z in Fig. 5) presents an important concentration of members. Moreover, if we consider only clusters with a higher membership probability (75 per cent, as noted in Fig. 6), and thus eliminating most of the clusters around the line of nodes in which the GB plane intersects the LGD plane, the GB is practically reduced to the Orion region, as if that were solely the cluster population of the GB.

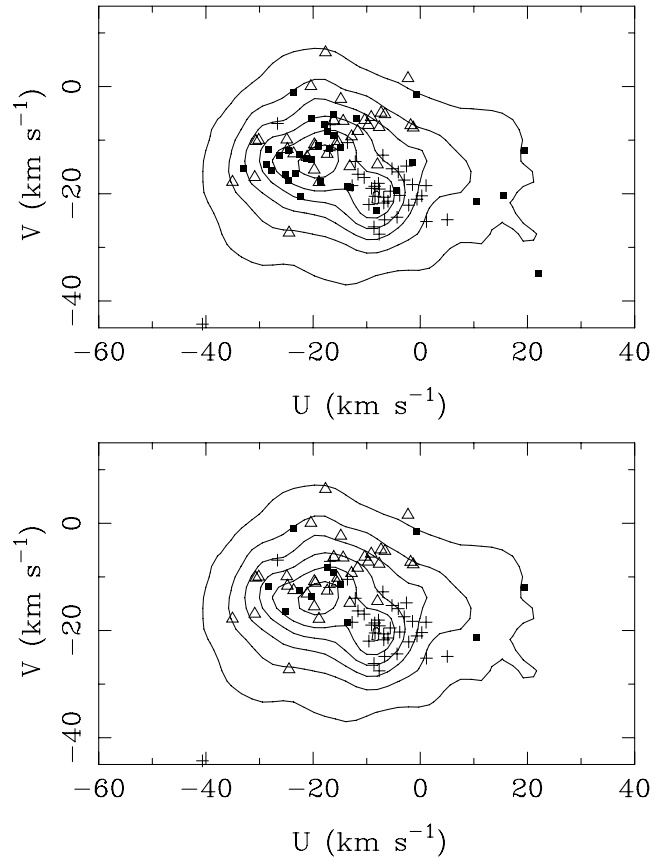


**Figure 6.** Same as Fig. 4, but including only GB clusters (filled circles) with a membership probability higher than 75 per cent. Other symbols are as in Fig. 3.

This analysis shows as an evident fact that while the Ori OB1 association is related to an important population of star clusters (be they transitory stellar condensations or gravitationally bound systems), the Sco OB2 complex does not include a single star cluster within its frontiers, yet it is extremely rich in massive OB stars and pre-main-sequence stars (Sartori et al. 2003). In other words, inside an apparently single star formation complex as the GB, and separated only by about 500 pc, we find two star-forming regions, well defined both spatially and kinematically, that present two clearly distinct modes of star formation. Orion is an OB association that presents a large number of stellar condensations seen as clusters, some of which appear to be gravitationally bound. On the other hand, Sco-Cen seems to be forming only isolated stars or loose groups that do not present the shape of a star cluster, and thus have not been detected as such. But, how does this fact translate to the velocity space?

### 3.2 Kinematic behaviour

It seems obvious that the different spatial position of the centroids of the two stellar groups analysed in Section 3.1 implies also a different location in the velocity space, specially in the  $V$  component due to their separation of almost 500 pc in the  $X$ -axis. The velocity centroids of these two associations, Ori OB1 and Sco OB2, as defined by the GB stars, are located at  $(-16.4, -9.5, -5.0)$  and  $(-8.0, -19.4, -6.0)$   $\text{km s}^{-1}$ , respectively. This can be clearly observed in Fig. 7 (top panel), where we have represented in the  $UV$  plane the isodensity lines of the GB as defined by massive stars, as well as the member stars of Ori OB1 and Sco OB2, according to

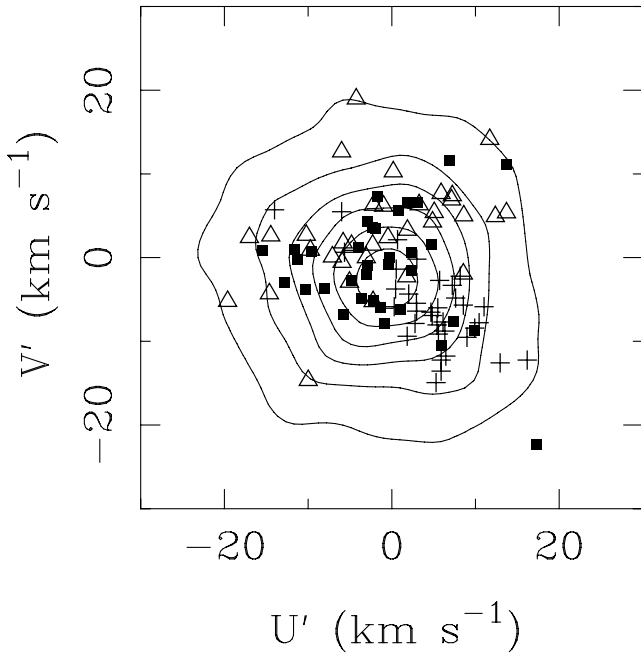


**Figure 7.** Velocity distribution of GB clusters (filled squares) against that of the GB stars (isodensity lines) from Paper I. Crosses and open triangles mark, respectively, the stars belonging to Sco OB2 and Ori OB1, according to de Zeeuw et al. (1999). Top panel represents the totality of GB clusters, whereas bottom panel displays only those classified as *probable* transient stellar condensations.

the criteria by de Zeeuw et al. (1999) explained in Section 3.1, and the GB star clusters. Once more, as it was expected, we observe a clear correlation between the clusters' distribution and the Orion velocity centroid, while the number of GB clusters associated with the Sco-Cen complex is merely marginal.

The difference is even more evident if we limit ourselves to the youngest clusters, which could be representatives of transient stellar condensations (Fig. 7, bottom panel). As we have commented above, the difference observed between the velocity centroids was predictable due to the differential Galactic rotation and to the fact that both groups were quite separated in space. However, we wonder if the rotation field in this region of the Galaxy can completely explain the kinematic behaviour of both the stars and the clusters in the GB.

Moreno, Alfaro & Franco (1999) analysed the velocity space of a sample of OB stars belonging to the GB, and found that the observed velocity field was not compatible with that obtained from star formation models, as well as with its dynamical evolution after the injection of momentum and energy in the primeval cloud. In other words, the residual velocities of the stars showed a highly negative vertex deviation (Filin 1957; Mihalas & Binney 1981) that could not be explained by the dynamical model, because after a time interval quite inferior than the age of the GB, the differential rotation prevailed over the movements originated by the 'ad hoc' energy and momentum injection, and gave place to a slightly positive vertex



**Figure 8.** Distribution of the residual velocities, corrected from solar motion and differential rotation, of GB clusters (filled squares) against that of the GB stars (isodensity lines) from Paper I. Crosses and open triangles mark, respectively, the stars belonging to Sco OB2 and Ori OB1, according to de Zeeuw et al. (1999).

deviation. Moreno et al. (1999) also found that if the stars belonging to the Sco-Cen association were eliminated, the vertex deviation became positive. That is, those models designed to explain the origin of the GB from supernovae explosions after a previous process of star formation, or from the impact of a high-velocity cloud on the Galactic disc (see Pöppel 1997, 2001; Sartori et al. 2003 for reviews about possibly formation mechanisms of the GB), are not able to explain the observed residual velocity field, unless the Sco-Cen stars are neglected.

Now we want to evaluate the residual velocity ellipsoid for the GB clusters; in order to do so, we have corrected the velocities from solar motion and differential rotation using the Oort constants:  $A = 16 \text{ km s}^{-1} \text{ kpc}^{-1}$  and  $B = -16 \text{ km s}^{-1} \text{ kpc}^{-1}$  (Paper II; Méndez et al. 2000). These cluster residual velocities are represented in Fig. 8, along with the residual velocities of OB stars, where different marks indicate those belonging to Ori OB1 and Sco OB2. Then, we calculate the longitude of the vertex,  $l_v$ , for both systems, the LGD and the GB cluster samples. The result, if we cut at GB membership probability of 50 per cent, is that  $l_v = 9.3 \pm 8.3$  for the LGD, and that  $l_v = -1.3 \pm 15.5$  for the GB. This is a value very far from the GB vertex deviation of  $l_v = -47^\circ$  found in Paper II for the OB stars belonging to the GB. This is undoubtedly caused by the absence of clusters in the Sco-Cen association, that was responsible for the large vertex deviation of the GB (Moreno et al. 1999; Paper II). Moreover, if we keep only GB clusters with a membership probability higher than 75 per cent, the vertex deviation of the system is  $l_v = 9.7 \pm 16.3$ , which is practically the same as that of the LGD.

The vertex estimation for the star cluster system in the GB thus gives us double information. First, from a kinematic point of view, the lack of star cluster associated with the Sco OB2 complex is demonstrated. Secondly, the difference in the velocity space be-

tween the Ori OB1 and Sco OB2 associations cannot be completely explained by the Galactic differential rotation.

As we may see in Fig. 8, the residual velocity distributions of these two associations show a different behaviour. While the stars belonging to Ori OB1 present a main velocity dispersion axis with a positive vertex, the Sco OB2 stars present a main axis that is almost perpendicular to the former, with a clearly negative vertex deviation. Therefore, as it had already been noted (Moreno et al. 1999; Paper II), the vertex deviation in the solar neighbourhood will depend on the selection of the sample. If most of the sample stars belong to Sco OB2, the vertex deviation will undoubtedly be negative. If we extend our sample farther away from the Sun in order to include Orion stars, we will reach some balance, and thus the vertex deviation will turn towards values closer to zero. The latter is precisely what we observe in the star cluster population associated with the GB: the lack of clusters within the Sco-Cen complex makes the vertex deviation close to zero.

### 3.3 Cluster complexes and scaled OB associations

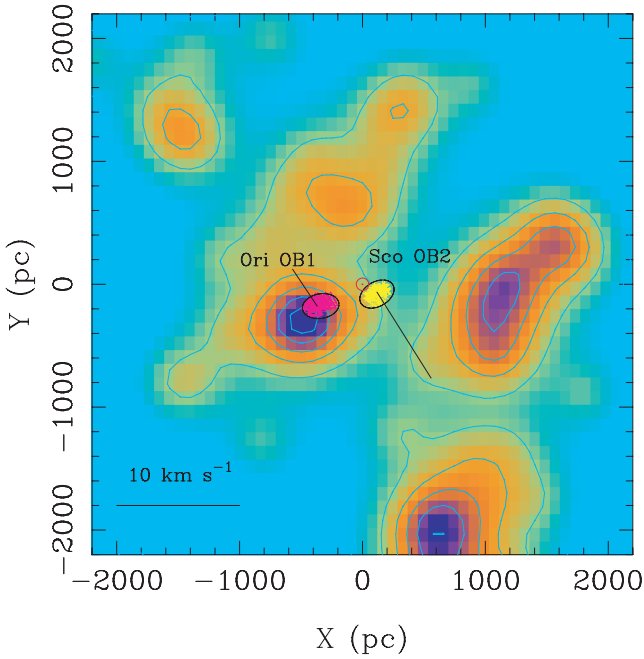
What we have observed when comparing the distribution of star clusters and OB stars in the GB is that the two great complexes that seem to define the north and south Galactic extremes of this large stellar structure show a different behaviour according to the scenery of hierarchical star formation. Ori OB1 shows a considerable portion of its stellar population as grouped and forming star clusters, half of them being probable bound clusters with ages larger than 10 Myr. This kind of stellar system has been detected, observed and analysed in both the Milky Way (Efremov & Sitnik 1988; Alfaro, Cabrera-Cano & Delgado 1991) and other galaxies (e.g. 30 Doradus in the LMC; Hunter et al. 1995; Walborn, Maíz-Apellániz & Barbá 2002); it not only contains a rich star cluster population, but also is normally immersed inside a stellar halo.

On the opposite side of the star-forming regions' concentration range, there can be found the scaled OB associations (SOBAs; e.g. Maíz-Apellániz 2001), of which NGC 604 in M33 is a good example. Although the star formation rate in NGC 604 is much higher than that observed in the Sco OB2 association, they both have in common their lack of star clusters. Thus, separated only by 500 pc, there are two OB associations that, apparently, have been born from molecular clouds under different ambient conditions, but that always have been considered as the fundamental parts of a single stellar system known as the GB. However, Ori OB1 seems to represent the stellar halo associated with a cluster complex, while Sco OB2 appears to be a clear example of an OB association, not related to cluster formation.

If we interpret this result in terms of a hierarchical star formation process (Elmegreen 2006, 2008), and considering that the age of the stars in both associations present similar ranges, we should consider either that the density maximum in the parental gas distribution of Sco OB2 was inferior than the density peaks in the clouds that formed Ori OB1, or that the tidal forces in the Sco-Cen region were intense enough to destroy any substructure observable as a star cluster, in an interval of time lesser than 10 Myr.

Any of these two possibilities requires some variations of the ambient physical conditions in spatial scales smaller than 500 pc, be they due to an external difference of pressure that caused higher density peaks in the Orion region, or/and due to local gravitational potentials or shear forces that shorten the lifetimes of the transient stellar condensations in the Sco-Cen region.





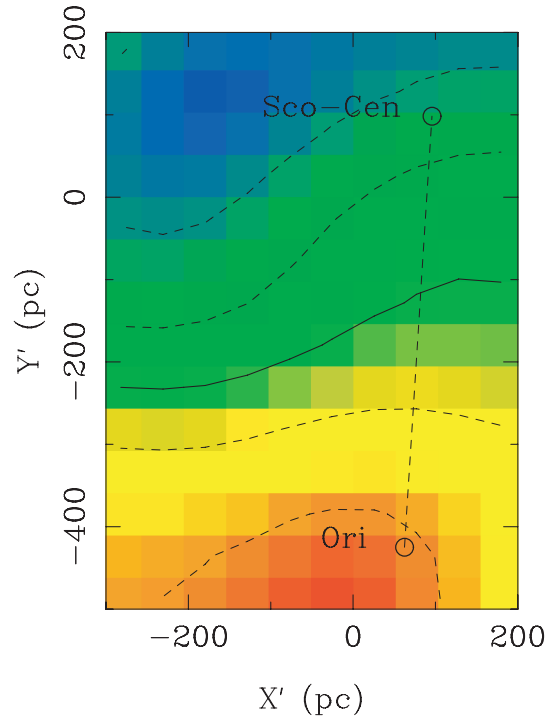
**Figure 9.** Density map for the star clusters younger than 10 Myr, within a box 4 kpc of side centred in the Sun (red dotted circle). The Ori OB1 and Sco OB2 associations have been superimposed on the map, along with their respective residual velocity vectors (black lines).

### 3.4 The Orion Arm

Since the pioneering work of Becker (1956), who traced the local spiral structure from the young star clusters distribution within a radius of 2 kpc, each time that a new catalogue of star clusters has been tailored, the corresponding map of this local spiral structure has been drawn again (e.g. Janes & Adler 1982). The results of doing this show a series of cluster groupings with typical sizes of the order of 1 kpc (Efremov & Sitnik 1988; Alfaro et al. 1991; Alfaro, Cabrera-Cano & Delgado 1992) that seem to follow three segments of spiral arms which have received the names of Carina–Sagittarius Arm, Perseus Arm and Orion Arm (also known as Local Arm). The inclusion of these optical segments within the general scheme of the spiral structure of the Galaxy is controversial, and depending on the spiral tracers and the analysis techniques employed, different solutions for the number of arms, their pitch angle or the velocity of the density pattern have been found (see Naoz & Shaviv 2007 and references therein).

In Fig. 9, we represent the density map for the star clusters younger than 10 Myr, inside a square of side length 4 kpc, centred in the Sun. The data have been extracted from the COCD catalogue. In the figure, we observe five main concentrations of young clusters that had previously been detected by other authors (e.g. Efremov & Sitnik 1988). In particular, the Orion and Cygnus complexes seem to align, delimiting the local optical segment of the Orion Arm. Superimposed over the young clusters density map, we have drawn the Ori OB1 and Sco OB2 associations. The lines that depart from the centroids of both associations represent their residual velocities, corrected from solar motion and differential rotation.

Ori OB1 appears associated with the density maximum of the Orion Arm defined by the young clusters; on the contrary, Sco OB2 is located in the inner rim of the Arm, close to the Sun where the cluster density is lower. If we consider that a higher density of clusters is representative of a higher ambient pressure in the original



**Figure 10.** Density distribution of the CFI in the GB plane. The open circles represent the centroids of the Ori OB1 and Sco OB2 associations. The solid contour line represents the CFI value 0.5, and the dashed contour lines are separated by a value of the CFI of 0.05.

gas, the relative position of these associations with respect to the main locus of the Orion spiral arm could explain their different content of star clusters.

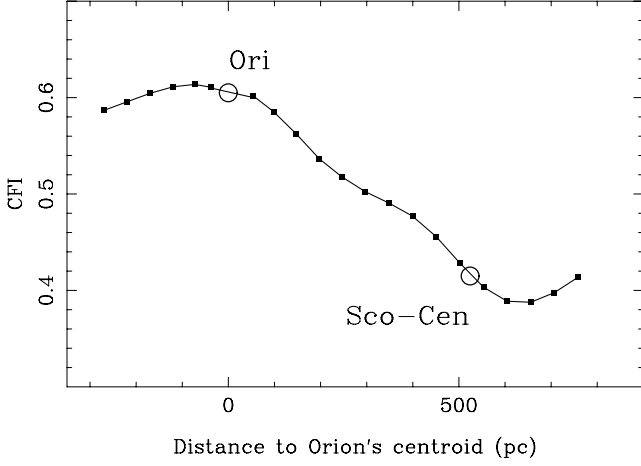
To better illustrate this, we define a Cluster Formation Index (CFI) that describes the relative content of clusters with respect to the OB stars that shape the GB:

$$\text{CFI} = \frac{\text{Cluster density}}{\text{Cluster density} + \text{OB star density}}. \quad (1)$$

The spatial densities of clusters and OB stars have been obtained by using Gaussian kernels in the GB plane ( $X'Y'$ ), with a  $\sigma = 200$  pc, and normalized so that the total sum of the density be equal to 1. Then, Fig. 10 shows the distribution in the GB plane of the CFI parameter in the region that contains both OB associations. A line that joins the centroids of both associations has been represented over the density map, and a cut along this line (Fig. 11) clearly shows a CFI gradient that ranges from 0.62 at the maximum near Orion to 0.39 in the vicinity of Scorpius Centaurus. Since the distance to the Ori OB1 centroid is a good estimator of the distance to the Orion spiral arm, this indicates that as we move away from the arm, the clustered star formation is less efficient.

### 3.5 The nature of the Gould Belt

The first detection of the GB was based on the fact that the brightest stars in the sky, especially in the Southern Hemisphere, present an eccentric position with respect to the great circle of the Milky Way (Herschel 1847; Gould 1879). That is, it was a mere morphological matter. Later studies determined the main stellar and gaseous components of the GB, and from the analysis of their spatial and kinematical properties, the shape, size and kinematics of the GB were estimated.

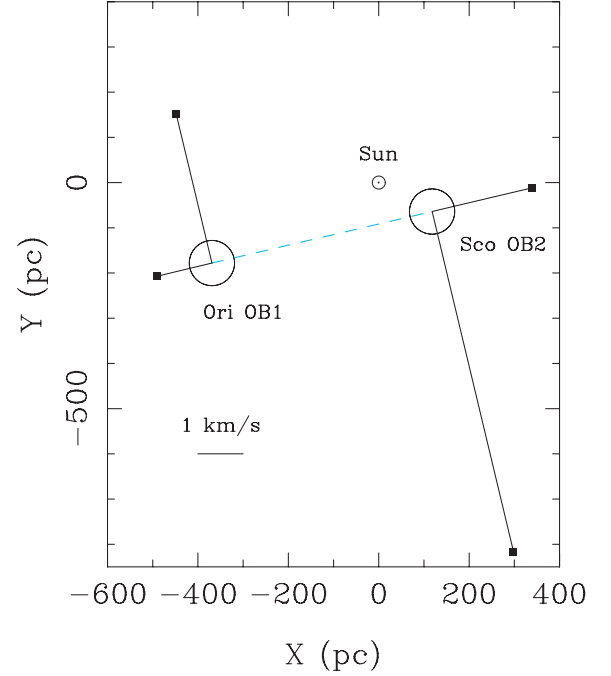


**Figure 11.** Value of the CFI along the line that joins the Ori OB1 and Sco OB2 centroids (from Fig. 10).

Although the main parameters that describe the GB, published along the past two decades, may differ depending on the authors (see Paper I; Pöppel 1997 and references therein), the general scheme of the GB is similar for most of them: we are facing a star formation complex with a disc-like structure and a radius of about 500 pc, whose kinematic behaviour is characterized by an expansion and a rotation with respect to an internal axis. In most of the works that contributed to the determination of these characteristics, the separation between the probable members of the GB and the stars belonging to the LGD has been performed. The analysis of the stellar component of both groups has led to the conclusion that the GB and the LGD are not only separated in the celestial sphere, but also present different kinematic properties as well.

In most of these analyses, it was necessary to perform a previous separation of the elements belonging to either group, and in all those cases such separation was mainly geometric (be it either in the  $l$ - $b$  plane, in the different  $XY$ ,  $XZ$ ,  $YZ$  projections, or via a three-dimensional analysis). In other words, the stars that are the brightest and the farthest from the Milky Way plane, and that belong to the Sco-Cen and Orion constellations, seem to have been the ones that opened the gates to this flood of studies, and also the ones that, still today, best define and delimit the geometry of the GB.

Then, we shall perform the following experiment: let us consider the GB as defined only by two points whose coordinates in the phase space are given by the spatial coordinates of the centroids of the Ori OB1 and Sco OB2 association, and their velocities by their respective central values of their residual velocities (Table 2). A schema of the geometry of the problem is shown in Fig. 12. Using these values, we determine the inclination ( $i$ ), the longitude of the ascending node ( $\Omega$ ), the expansion velocity along the line that joins the spatial centroids ( $\rho_0$ ) and the rotation velocity with respect to a point situated along the main axis of the system ( $\omega_0$ ), we reach the results listed in Table 3, where we also present the range of values found in the literature for the different variables in the GB.



**Figure 12.** ‘Reduced GB’, showing the system as a schematic composition of the centroids of Ori OB1 and Sco OB2 and their residual velocities (listed in Table 2).

**Table 3.** Comparison of the ‘reduced GB’ parameters with their range of values found in the literature.

	$i$ ( $^\circ$ )	$\Omega$ ( $^\circ$ )	$\rho_0$ ( $\text{km s}^{-1} \text{ kpc}^{-1}$ )	$\omega_0$ ( $\text{km s}^{-1} \text{ kpc}^{-1}$ )
Reduced GB	18	283	7	22
Literature	14–27	271–290	0–29	12–37

These values are all very close to those of the fundamental parameters of the GB that we can find in the literature (e.g. Paper I; Fresneau et al. 1996; Lindblad 2000, among others). But in view of Fig. 9, where we represent the position and the residual velocities of both associations with respect to the Orion Arm, we may ask ourselves if it makes any sense at all talking about a gas and stars system with spatial and kinematic coherence that may be described with some expansion and rotation velocities, or perhaps we should definitely drop the traditional hypothesis of a single, common origin for all the features of the GB, and begin to look at it as a hazardous alignment – from our point of view – of at least two of the many clumps in the hierarchical structure of the Local Arm, with different densities and star formation histories. In this sense, the GB would be simply the projection over the sky of the recent star formation in the clouds close to the Sun but located far away from the fundamental Galactic plane. Therefore, it is morphologically distinct, but does

**Table 2.** Centroids of the OB associations Ori OB1, and Sco OB2, in phase space.

Name	$X$ (pc)	$Y$ (pc)	$Z$ (pc)	$U'$ ( $\text{km s}^{-1}$ )	$V'$ ( $\text{km s}^{-1}$ )	$W'$ ( $\text{km s}^{-1}$ )
Ori OB1	$-369 \pm 7$	$-178 \pm 6$	$-132 \pm 4$	$-2 \pm 1$	$3 \pm 1$	$1 \pm 0.6$
Sco OB2	$118 \pm 7$	$-64 \pm 6$	$29 \pm 4$	$4 \pm 1$	$-7 \pm 1$	$0 \pm 0.5$

not necessarily represent a physical system with unique properties, different from the star formation within the Orion Arm.

We believe that the distinct properties of these two clumps – their kinematic behaviour, cluster content and height over the Galactic plane – may be explained by the internal dynamics of the Galactic disc. Further research should consider, as possible mechanisms involved in the development of such a structure, the passing of a density wave in a magnetized medium or the presence of a long bar generating resonances in the external parts of the disc (Gardner, Innanen & Flynn 2008).

Thus, if we could observe our Galaxy from an external position several megaparsec away, what would stand out in the solar neighbourhood? Undoubtedly, the complex of blue star clusters inside the Local Arm, where the Ori OB1 and Sco OB2 associations, that form the main structure of the GB, would just be accessory elements related to the nucleus of the complex and its periphery, respectively.

## 4 CONCLUSIONS

From a sample of star clusters younger than 100 Myr, and located within a radius of 1 kpc around the Sun, we have analysed the spatial and kinematical structure of the GB. The comparative analysis between the populations of stellar clusters and OB stars in the GB indicates that the Ori OB1 and Sco OB2 associations present a significant difference in the number of clusters related to them, and that this difference is even more conspicuous in the case of the transient stellar condensations. While Ori OB1 can be characterized as the stellar population associated with the core of a star cluster complex (e.g. 30 Doradus), the Sco-Cen complex apparently shows a star formation mode where the generation of isolated stars is dominant.

In the light of this scenario of hierarchical star formation, this difference in the content of stellar clusters must have been caused by different physical conditions of the primeval clouds and/or inhomogeneities of the gravitational potential. The main physical characteristics of these two large associations are as follows.

- (i) Different height over the Galactic plane.
- (ii) Different content of stellar clusters.
- (iii) Different residual velocity vectors.

All these differences can be explained, at least qualitatively, by the different position of these two associations with respect to the main loci of the young stellar clusters that define the Orion Arm. According to this scenario, the GB can be considered as a partial and biased vision of a much larger scale process of stellar formation, which is currently visible as a star cluster complex in the region of the Orion Arm that is closest to us.

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