

The luminosity function of the Virgo Cluster from $M_B = -22$ to -11

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ABSTRACT

We measure the galaxy luminosity function (LF) for the Virgo Cluster between blue magnitudes $M_B = -22$ and -11 from wide-field charge-coupled device (CCD) imaging data. The LF is only gradually rising for $-22 < M_B < -16$. Between $M_B = -16$ and -14 it rises steeply, with a logarithmic slope of $\alpha \sim -1.6$. Fainter than $M_B = -14$, the LF flattens again. This LF is shallower (although turning up at brighter absolute magnitudes) than the R -band LF measured recently by Phillipps et al., who found $\alpha \sim -2.2$ to be fainter than $M_R = -13$. It is similar, however, to the LF determined from the Virgo Cluster Catalog by Sandage et al. A few faint galaxies are found that Sandage et al. missed because their surface brightness threshold for detection was too high, but these do not dominate the luminosity function at any magnitude. Most of the faint galaxies we find are dwarf elliptical, alternatively called dwarf spheroidal, galaxies. The most important potential source of systematic error is that we may have rejected some high surface brightness galaxies from the cluster sample because we think that they are background galaxies. This is quite different from what has conventionally been regarded as the most serious source of systematic error in this kind of study: that we are missing many *low* surface brightness galaxies because they are never visible above the sky.

There are ~ 2.5 times more dwarfs per giant galaxy in Virgo than in the Ursa Major Cluster, a diffuse group of ~ 80 spiral galaxies at the same distance as Virgo, or the Local Group. The Virgo and Ursa Major Cluster LFs are inconsistent with each other at a high level of significance. These results add weight to the hypothesis that is developing that dwarf galaxies are more common relative to giant galaxies in dense environments than diffuse ones. Both LFs are highly inconsistent with cold dark matter theory, which has been so successful at reproducing observations on large scales. Possible theoretical explanations for this discrepancy, and for the detailed shape of the Virgo Cluster LF, are investigated.

Key words: galaxies: clusters: individual: Virgo – galaxies: luminosity function, mass function.

1 INTRODUCTION

The Virgo Cluster is one of only two elliptical-rich clusters within 25 Mpc, the other being the Fornax Cluster. It is near enough that even low-luminosity galaxies appear quite big on the sky and can be identified and studied in some detail.

1 arcsec corresponds to ~ 0.1 kpc in Virgo, which is comparable to the scalelengths of low-luminosity galaxies (Binggeli & Cameron 1991). What this means is that if we see a $B > 18$ galaxy in the Virgo Cluster with a size of several arcsec, it is much more likely to be a low-luminosity cluster member than a high-luminosity background galaxy, since intrinsically lower-luminosity galaxies have lower surface brightnesses and consequently larger

scalelengths (Binggeli 1994). A comparison of the joint apparent magnitude–surface brightness distribution between the Virgo Cluster and blank sky fields therefore gives some indication of the low-luminosity galaxy content of the Virgo Cluster.

Sandage, Binggeli & Tammann (1985) measured the luminosity function of the Virgo Cluster using the Virgo Cluster Catalog (VCC; Binggeli, Sandage & Tammann 1985) and found a luminosity function $\phi(L)$ that was gently rising at the faint end: $\alpha \sim -1.35$, where $\phi(L) \propto L^\alpha$. They detected galaxies with total B magnitudes $B_T \sim 20$ (approximately $M_B = -11$), although they estimated that their completeness limit was two magnitudes brighter than this: at $B_T > 18$, significant numbers of dwarfs have surface brightnesses so low and scalelengths so large that they were not detectable above the night sky in the photographic images used to compile the VCC. Following this work, Impey, Bothun & Malin

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(1988) discovered a number of low surface brightness galaxies with $B_T < 20$ missing from the VCC. They suggested that incompleteness at the faint end of the VCC might be very severe and that the LF could be as steep as $\alpha = -1.7$. In a more recent development, Phillipps et al. (1998a) present evidence for large numbers of low-luminosity galaxies in the Virgo Cluster. They measure $\alpha \sim -2.2$ fainter than an absolute R magnitude of $M_R = -13$. If the LF was this steep, Sandage et al. would have missed the vast majority of cluster members at the faint end, even after imposing their completeness corrections.

All of these results suggest that the LF is very probably steeper in the Virgo Cluster than in diffuse spiral-rich groups and clusters. In the Local Group, where very faint absolute magnitudes ($M_B \sim -8$) can be reached, the faint-end slope is $\alpha = -1.1$ (van den Bergh 1992, 2000). In the Ursa Major Cluster, a diffuse spiral-rich group at a similar distance to the Virgo Cluster, less faint absolute magnitudes can be reached, but large enough numbers of galaxies are present that a statistically robust LF can be computed. Here $\alpha = -1.1$ as well (Trentham, Tully & Verheijen 2001a). Evidence is therefore accumulating that low-luminosity galaxies are very much more numerous per luminous galaxy in dense environments than in diffuse ones.

Values of $\alpha \sim -2$ are of particular interest since this is the logarithmic slope of the low-mass galaxy mass function predicted by theory if the primordial fluctuation spectrum is a power law with index $n = -2$, as is appropriate for cold dark matter (Press & Schechter 1974; White & Rees 1978; Lee & Shandarin 1999; Klypin et al. 1999). If this value of α is appropriate for the Virgo Cluster, this would suggest that in this environment the efficiency of star formation in small galaxies does not depend on the galaxy mass, assuming cold dark matter theory [the problem of reproducing the shallow LF slope $\alpha \sim -1$ in the diffuse environments still remains (Moore et al. 1999; Klypin et al. 1999)].

We now present the results of a survey of 25 deg^2 of the Virgo Cluster (approximately one-tenth of the total area of the cluster) observed through the B filter using the Wide Field Camera on the Isaac Newton Telescope on La Palma, taken as part of the INT Wide Field Survey. The intention is to obtain a reasonably complete (at least to the surface brightness levels defined in the studies mentioned above) inventory of Virgo Cluster members. We can then construct a luminosity function from $M_B = -22$ (the brightest galaxy in our sample was M87 with $M_B = -21.5$) down to $M_B = -11$. This will permit us to address the following questions.

(i) What is the value of α , and does this vary significantly with absolute magnitude, i.e. over what magnitude range can we approximate the LF by a power law? How sensitive is the answer to our ability to recognize cluster members based on surface brightness: could we be missing many high surface brightness members because we think that they are background galaxies or many low surface brightness galaxies the contrast of which against the sky is too low to allow us to identify them?

(ii) What are the morphologies of the faintest galaxies? Most previous work (Sandage et al. 1985; Phillipps et al. 1998a) suggests that they are dwarf elliptical (alternatively called dwarf spheroidal) galaxies. This would imply that low surface brightness irregular star-forming galaxies do not contribute significantly to the LF at the faint end as they do in the Ursa Major Cluster (Trentham et al. 2001a).

(iii) Do the results depend on the colour of the galaxies and the filter used? We, like Sandage et al., are using a B filter. Phillipps

et al. used a red R filter. How much of the very substantial excess of galaxies they found could be due to this, given that dSph/dE galaxies tend to be red (Caldwell 1983)?

(iv) Are there substantial numbers of very low surface brightness (VLSB) galaxies, as seen in the Fornax cluster by Kambas et al. (2000)? In particular, as we approach the limiting surface brightness at which we can detect objects in our data, do we find more and more VLSB galaxies, and what is their contribution to the total LF? We will need to quantify this in order to address the questions posed in (i) above.

(v) How does the galaxy luminosity function depend on environment? We will compare our Virgo Cluster luminosity function with the B -band luminosity functions of the Local Group, the Ursa Major Cluster, and the rich Coma Cluster at a distance of 90 Mpc. We will then have measurements of the LF in four very different environments. In the case of the Local Group, the LF has large uncertainties owing to Poisson statistics. In the Coma Cluster the LF has large uncertainties because of the need to do a background subtraction. It is well known (Dressler 1980) that the morphologies of galaxies depend on environment, specifically on the galaxy density. Our results will indicate whether or not the LF does too, over a large magnitude range.

2 OBSERVATIONS

The data used here were taken on various observing runs during Spring 1999 and Spring 2000 as part of the INT Wide Field Survey (WFS; <http://www.ast.cam.ac.uk/~wfcSUR>; McMahon et al. 2001). This is a digital survey covering $\sim 100 \text{ deg}^2$ of sky carried out using the Wide Field Camera [a mosaic of four $4K \times 2K$ EEV charge-coupled devices (CCDs), pixel scale $0.33 \text{ arcsec pixel}^{-1}$, field of view 0.29 deg^{-2} ; <http://www.ast.cam.ac.uk/~wfcSUR/ccd.html>] on the 2.5-m Isaac Newton Telescope on La Palma.

A total of 113 fields in the Virgo Cluster were selected (see Fig. 1), comprising 24.9 deg^2 taking into account overlaps between

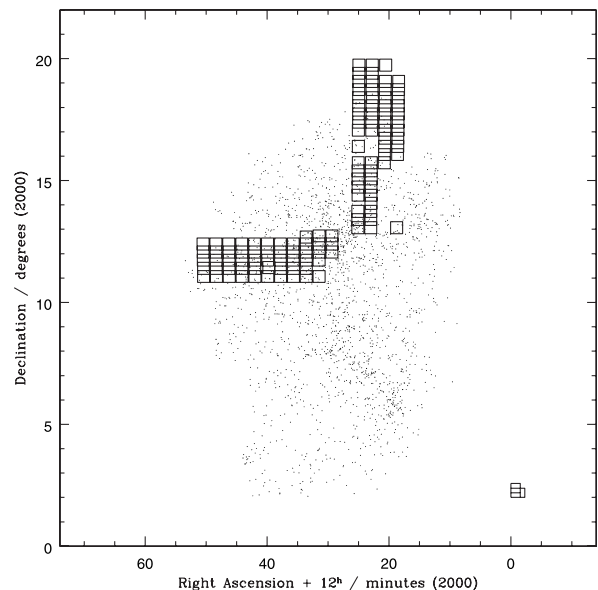


Figure 1. The region of the Virgo Cluster studied. The open boxes represent the 113 fields observed; the box size corresponds approximately to the field of view of the INT Wide Field Camera. The dots represent the locations of the 2096 Virgo Cluster Catalog galaxies (Binggeli et al. 1985). The exact shape of the INT Wide Field Camera, indicating the positions of the four CCDs, is shown in the bottom right-hand corner.

the images, which were typically one chip or 1/4 of the field of view. All fields were imaged through a B (effective wavelength $\lambda_0 = 0.44 \mu\text{m}$) filter for 750 s under photometric conditions. A total of 31 deg^2 from the WFS imaging survey of the North Galactic Cap (NGC), taken in photometric conditions with a B filter, were also studied, these being used as offset fields to estimate the background contamination in the Virgo sample. The exposure times in B (750 s) were the same for the Virgo and NGC data.

The data were pre-processed and reduced via the WFS pipeline (<http://www.ast.cam.ac.uk/~wfsur/pipeline.html>; Irwin & Lewis 2001) and photometric calibration was obtained from observations of several (5–10 per night) standard stars. The photometric zero points were always accurate to 2 per cent; uncertainties in the zero point are not a major source of uncertainty in the galaxy magnitudes that we derive. The median seeing was 1.95 arcsec for the Virgo data and 1.26 arcsec for the NGC data (since we will be searching for galaxies with large sizes and low surface brightnesses, the relatively poor seeing of the Virgo data is not a major concern). For the Virgo data, the median 3σ point source limiting magnitude was $B = 25$.

Of the 113 fields we studied, nine of the fields were also imaged through a Z ($\lambda_0 = 0.90 \mu\text{m}$) filter for either 600 or 1200 s under photometric conditions. These data were used to compute $B - Z$ colours for a small number of galaxies in the B sample. The median seeing in the Z data set was 1.53-arcsec FWHM.

3 GALAXY IDENTIFICATION AND SELECTION

Initially, an image detection algorithm developed by Irwin (1985, 1996) was run on the images. A source was defined to be five contiguous pixels with a flux signal-to-noise ratio of 1.5 above the sky. Generated parameters include the position, intensity and shape of each detected object. We used only the position information. Aperture magnitudes were subsequently calculated for all the sources using the IRAF APPHOT task with a local sky value measured for each object (typically the mode within an annulus of inner and outer radii 0.5 and 1.0 arcmin centred on the object).

In our data set we had $\sim 10^6$ galaxies ($\sim 2500 \text{ chip}^{-1}$) with B magnitudes within a 6-arcsec aperture brighter than $B(6) = 23$ [here, as throughout the paper we use the notation $B(r)$ to mean the magnitude within an aperture of radius r arcsec]. In this section we describe how we identified which ones are likely to be Virgo Cluster members and assess how much confidence we should have in these identifications.

Many of the brightest galaxies [typically those with $B(6) < 18$] had spectroscopic velocity measurements available. For such objects we can identify with 100 per cent confidence which are members and which are not. We follow Binggeli et al. (1985) in assigning cluster membership to all galaxies with heliocentric velocities between -700 and 2700 km s^{-1} .

For the other galaxies, the situation is more complex. As outlined in Section 1, we expect dwarf galaxies in the cluster to have lower surface brightnesses (and consequently more extended light profiles) than background galaxies of the same apparent magnitudes. Morphologically we expect them to be either smooth (dwarf elliptical/spheroidal galaxies) or lumpy (dwarf irregulars) but *not* to have appreciable spiral structure or a distinctive bulge + disc morphology, both of which are characteristic of luminous giant galaxies (Binney & Merrifield 1997), which here would be background galaxies.

We used both surface brightness and morphology to assess

membership as follows. First, we compared the light concentrations of galaxies in the Virgo and background samples. For each galaxy we computed (following Trentham et al. 2001a) an inner concentration parameter (ICP) and an outer concentration parameter (OCP), defined as

$$\text{ICP} = B(4.4) - B(2.2)$$

$$\text{OCP} = B(12) - B(6).$$

The results are presented in Fig. 2.

There is a clear excess of objects with $B(4.4) > 18$ and $\text{ICP} \sim -1.3$ in the Virgo fields that are not seen in the background fields. These are presumably Virgo Cluster members. They have the correct sizes for dwarf galaxies seen at the distance of the Virgo Cluster since they are close to the dashed lines in the upper left-hand panel of Fig. 2. The scatter around the dashed lines is huge, however, particularly in the lower panels. The dependence on surface brightness of the location of galaxies in the panels in Fig. 2 is complex. In general, lower-surface brightness galaxies have more extended light profiles yet it is not *always* true that lower surface brightness galaxies always have more negative concentration parameters. In the lower panels very low surface brightness galaxies have very small differences between their 6- and 12-arcsec aperture fluxes since so much of the galaxy falls below the sky; if the galaxy is faint (as for galaxies at the right end of the dashed lines), the only difference between $B(6)$ and $B(12)$ is caused by sky noise. This could lead to an OCP value that could be as high as ∞ ! Hence there is huge scatter at the faint end in the lower two panels. What all this means is that we must consider the $B(6)$ and concentration parameter values in conjunction with each other when assessing membership, i.e. in effect we must consider the entire light profile.

We therefore define a quantity

$$P = \frac{v[\text{OCP}, B(6)] - b[\text{OCP}, B(6)]/1.25}{v[\text{OCP}, B(6)]}, \quad (1)$$

where $v[\text{OCP}', B(6)']$ is the number of objects in the Virgo sample with 6-arcsec aperture magnitudes within 0.5 mag of $B(6)'$ and an outer concentration parameter within 0.05 mag of OCP' and an ICP value lower than -0.9 mag and $b[\text{OCP}', B(6)']$ is the equivalent number for the background fields. The factor of 1.25 comes from the relative area samples by our Virgo and background surveys. Defined in this way P is then an estimate of the probability of whether or not a particular galaxy is a cluster member.

For cluster members with velocity measurements, $P \sim 1$ so we can have some confidence in using P values to assess membership. However, some grand-design spirals, which are clearly background (these have very negative OCPs caused by considerable amounts of star formation far from the galaxy centres) also have $P \sim 1$. Hence using P values as a direct measure of membership is not completely safe. Other (related) reasons why we might be suspicious are as follows.

(i) The scatter in Fig. 2 around the dashed lines is large (see the previous discussion), so on an object-by-object basis, the uncertainty in P is large.

(ii) Our background sample (even though it comprises $\sim 30 \text{ deg}^2$ of data) is not sufficiently large to compensate for shot noise and Poisson errors in $b[\text{OCP}', B(6)']$ in regions where the OCP values are large. For example, in the lower right-hand panel of Fig. 2, there are only three galaxies with an $\text{OCP} < -1.1$. These are all low surface brightness field galaxies.

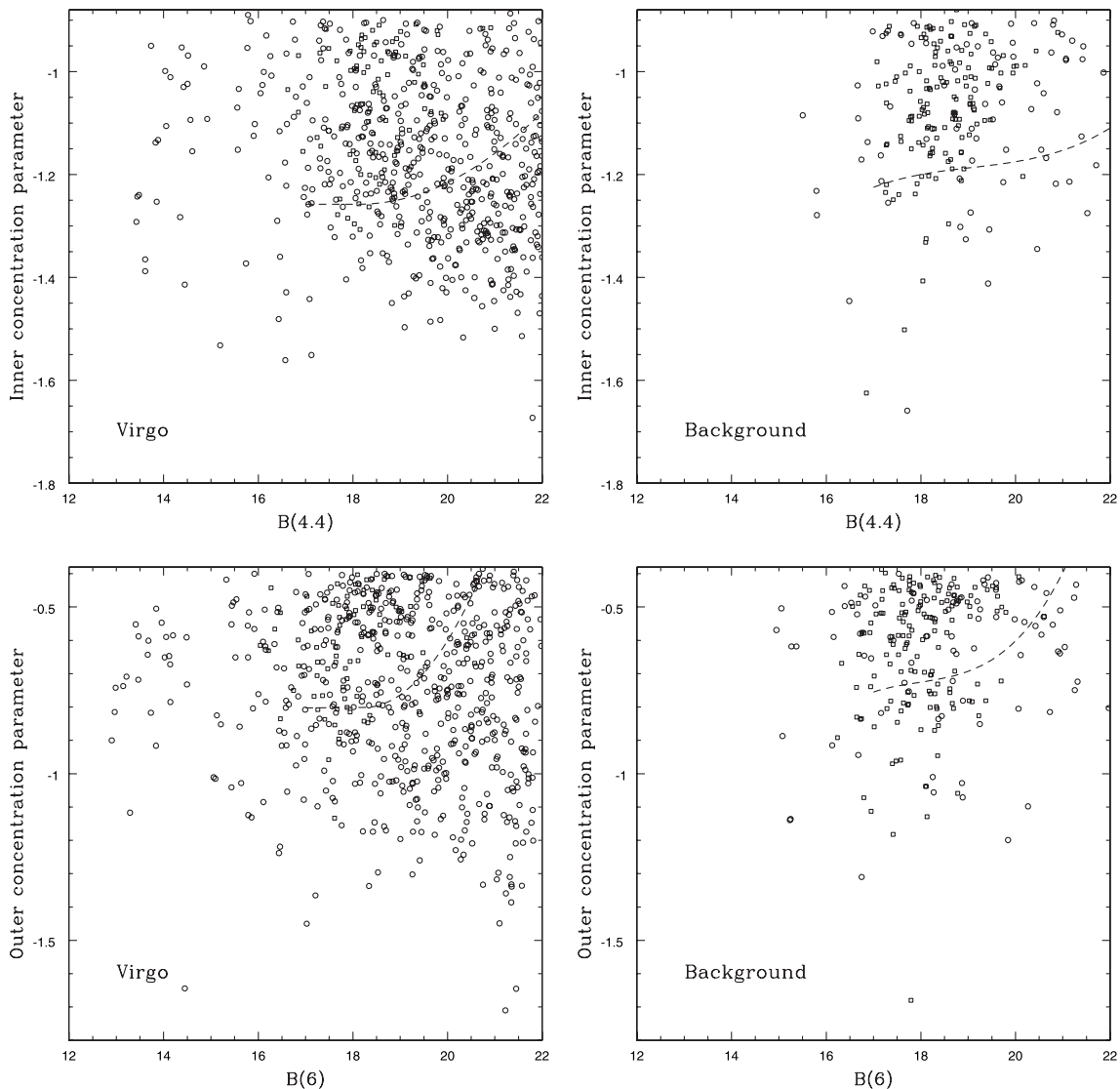


Figure 2. Inner and outer concentration parameters as a function of aperture magnitude for all objects we detected in both the Virgo and NGC (background) data that did not have close companions projected within 12 arcsec. Objects indicated with a square (as opposed to a circle) showed either clear spiral structure or were highly flattened discs, usually with central bulges; these are almost certainly background objects if seen in the Virgo fields. The dashed lines represent simulated typical dwarfs from fig. 1 of Binggeli (1994; surface brightnesses μ_B in mag arcsec $^{-2}$ and absolute blue magnitudes M_B are related by the approximation $\mu_B \approx 33 + 0.59M_B$) convolved with the seeing and ‘observed’ in the presence of noise appropriate to the data. Exponential profiles and axial ratios of 1 were assumed for the simulated dwarfs.

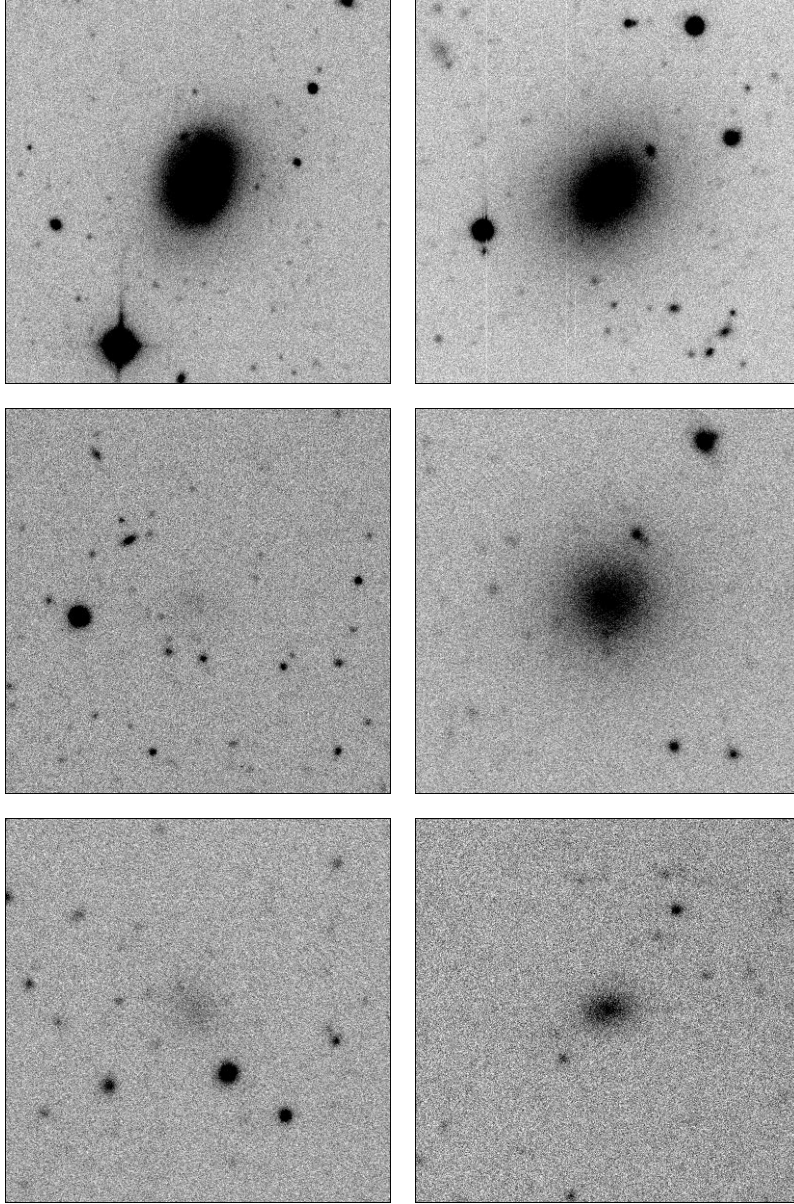
(iii) The field-to-field variance of these low surface brightness field galaxies is determined by large-scale structure at very low redshift and is consequently likely to be large. This makes the problem highlighted above in (ii) more severe. For example, suppose for some combination of $B(6)$ and OCP, we have 10 Virgo field galaxies with $ICP < -0.9$ mag and three background field galaxies. We would compute $P = (10 - 3/1.25)/10 = 0.81$. We would therefore think that most of these galaxies are cluster members. However, suppose large-scale structure at low redshift caused the field-to-field variance of these low surface brightness galaxies to be a factor of 4. Then the appropriate number of background galaxies with the relevant combination of parameters in the Virgo fields could now be 12, not three. So P should have been $(10 - 12/1.25)/10 = 0.04$, which is negligible. With this value of P , we would now think that most of these galaxies were *not* cluster members. In summary, until we have a characterization of

the field-to-field variance of field galaxies *as a function of surface brightness*, values of P must be treated with caution.

So instead of relying solely on P values, we further inspected each galaxy individually and made a judgement concerning the possibility of membership. We discarded galaxies showing obvious spiral structure and flat disc galaxies with a central bulge, even if their P values were high, since these are almost certainly background giant galaxies (~ 15 per cent of the galaxies that we would have identified as possible members in the absence of any morphological information beyond averaged light profiles were rejected on these grounds). For each galaxy that did not have a velocity measurement that we thought might be a member we rated it (as was done in Trentham et al. 2001a) ‘1’ or ‘2’ depending on whether we thought it was a member with a high or moderate degree of confidence. The rating scheme is summarized in Table 1.

Table 1. Rating scheme for Virgo galaxies.

Rating	Meaning	Comments
0	Confirmed member	Velocity measurements exist heliocentric velocity between -700 and 2700 km s^{-1}
1	Probable member	Velocity measurements do not exist; high P ; very low surface brightness
2	Possible member	Velocity measurements do not exist; moderate P ; low surface brightness, but not lower than the lowest surface brightness background galaxies


Figure 3. Sample galaxies from the catalogue, having $B(6)$ values of 17–22 in 1 mag units, clockwise from the top left. All images are square, 132 arcsec on a side.

In making this analysis, we relied for the main part on P values, but also on morphology, particularly in regions of parameter space that were poorly sampled in the background fields [see point (ii) above]. Most objects rated ‘2’ were galaxies of moderate surface brightness where there was some

overlap in parameter space between the background and Virgo; such galaxies were not rated ‘1’ because of point (iii) above.

There were a few additional complications resulting from this approach that required attention.

Table 2. The sample.

(1) ID	(2) Name	(3) Type	(4) α (J2000)	(5) δ (J2000)	(6) V_h (km s ⁻¹)	(7) $B(6)$	(8) ICP	(9) OCP	(10) B_T	(11) P	(12) Class	(13) M_B
1	M87 (NGC 4486)	E	12 30 49.46	12 23 28.8	1307	13.29	-1.25	-1.12	9.59	1.00	0	-21.64
2	M60 (NGC 4649)	E	12 43 38.02	11 33 11.2	1117	14.45	-1.51	-1.56	9.81	1.00	0	-21.45
3	M86 (NGC 4406)	E	12 26 11.64	12 56 56.6	-244	15.64	-1.53	-1.64	9.83	1.00	0	-21.45
4	M85 (NGC 4382)	S0	12 25 23.43	18 11 26.7	729	12.99	-1.24	-0.74	10.00	1.00	0	-21.28
5	M84 (NGC 4374)	E	12 25 03.76	12 53 13.4	1060	12.97	-1.24	-0.82	10.09	1.00	0	-21.23
6	M100 (NGC 4321)	Sbc	12 22 54.74	15 49 20.5	1571	13.84	-1.41	-0.92	10.05	1.00	0	-21.21
7	M58 (NGC 4579)	Sb	12 37 43.56	11 49 05.6	1519	13.47	-1.13	-0.72	10.48	1.00	0	-20.85
8	M59 (NGC 4621)	E	12 42 02.25	11 38 48.8	410	13.21	-1.39	-0.71	10.57	1.00	0	-20.72
9	M89 (NGC 4552)	E	12 35 39.83	12 33 23.4	340	13.14	-1.37	-0.74	10.73	1.00	0	-20.60
10	NGC 4293	S0/a	12 21 12.84	18 22 58.6	893	15.09	-1.03	-1.02	11.26	1.00	0	-20.06
11	NGC 4568	Sbc	12 36 34.30	11 14 20.8	2255	15.43	-1.13	-1.04	11.68	1.00	0	-19.61
12	NGC 4394	Sb	12 25 55.67	18 12 50.4	922	14.20	-1.09	-0.59	11.73	1.00	0	-19.55
13	NGC 4567	Sbc	12 36 32.71	11 15 28.4	2274	17.21	-1.26	-1.37	11.79	1.00	0	-19.40
14	NGC 4647	Sc	12 43 32.31	11 34 54.7	1422	15.06	-1.15	-1.01	11.94	1.00	0	-19.32
15	NGC 4503	S0	12 32 06.26	11 10 35.5	1342	14.02	-0.95	-0.65	12.05	1.00	0	-19.32
16	NGC 4694	S0	12 48 15.14	10 59 00.1	1175	14.14	-0.97	-0.67	12.06	1.00	0	-19.26
17	NGC 4564	E	12 36 27.02	11 26 21.7	1142	13.68	-1.00	-0.60	12.05	1.00	0	-19.25
18	NGC 4567	Sab	12 36 33.29	11 15 45.0	2274	16.44	-1.23	-1.24	12.06	1.00	0	-19.23
19	NGC 4660	E	12 44 32.01	11 11 25.9	1083	13.47	-1.14	-0.59	12.16	1.00	0	-19.13
20	NGC 4638	S0	12 42 47.38	11 26 32.4	1164	13.66	-1.11	-0.64	12.13	1.00	0	-19.13
21	NGC 4478	E	12 30 17.46	12 19 43.2	1349	13.73	-1.01	-0.82	12.36	1.00	0	-18.90
22	NGC 4486A	E	12 30 57.76	12 16 14.5	450	13.41	-0.95	-0.55	12.44	1.00	0	-18.81
23	NGC 4550	S0	12 35 30.62	12 13 15.0	381	14.14	-1.16	-0.79	12.56	1.00	0	-18.76
24	NGC 4312	Sab	12 22 31.39	15 32 15.9	153	15.86	-1.29	-1.13	12.53	1.00	0	-18.74
25	NGC 4402	Sb	12 26 07.01	13 06 48.4	232	17.02	-1.32	-1.45	12.55	1.00	0	-18.73
26	NGC 4379	S0	12 25 14.78	15 36 26.9	1069	14.11	-1.02	-0.65	12.63	1.00	0	-18.62
27	NGC 4606	Sa	12 40 57.48	11 54 42.1	1664	15.21	-1.37	-0.85	12.67	1.00	0	-18.62
28	NGC 4383	S0	12 25 25.53	16 28 12.8	1710	13.84	-0.83	-0.51	12.67	1.00	0	-18.58
29	NGC 4377	S0	12 25 12.46	14 45 42.0	1371	13.96	-1.28	-0.55	12.76	1.00	0	-18.55
30	NGC 4733	S0/a	12 51 06.80	10 54 43.6	908	15.12	-1.07	-0.83	12.70	1.00	0	-18.54
31	NGC 4528	S0	12 34 06.10	11 19 16.9	1374	14.05	-1.03	-0.59	12.97	1.00	0	-18.38
32	NGC 4551	S	12 35 37.98	12 15 50.7	1172	14.49	-1.09	-0.73	12.97	1.00	0	-18.35
33	NGC 4476	S0	12 29 59.14	12 20 55.6	1978	14.48	-0.99	-0.59	13.01	1.00	0	-18.26
34	NGC 4396	Sd	12 25 59.20	15 40 14.9	-128	16.46	-1.44	-1.22	13.06	1.00	0	-18.20
35	NGC 4497	S0/a	12 31 32.56	11 37 29.3	1123	15.61	-1.04	-0.86	13.19	1.00	0	-18.14
36	NGC 4344	BCD	12 23 37.76	17 32 27.0	1142	15.79	-1.48	-1.12	13.34	1.00	0	-17.91
37	NGC 4336	S0/a	12 23 29.85	19 25 36.9	1031	15.64	-1.03	-1.03	13.48	1.00	0	-17.84
38	NGC 4607	Sb	12 41 12.26	11 53 06.5	2257	16.84	-1.19	-1.06	13.75	1.00	0	-17.54
39	IC 3475	dI	12 32 40.65	12 46 10.0	2583	18.34	-1.28	-1.34	13.82	1.00	0	-17.44
40	NGC 4328	S0	12 23 20.05	15 49 13.5	499	16.49	-1.28	-0.92	14.04	1.00	0	-17.22
41	IC 3718	dI	12 44 45.39	12 21 03.2	849	16.61	-1.23	-1.05	14.07	1.00	0	-17.21
42	IC 3499	S0/a	12 33 45.02	10 59 45.0	1212	15.52	-1.10	-0.65	14.12	1.00	0	-17.17
43	IC 3470	dE,N	12 32 23.42	11 15 47.0	1500	16.00	-1.12	-0.76	14.29	1.00	0	-17.09
44	NGC 4641	S0	12 43 07.64	12 03 03.4	2017	15.80	-1.00	-0.52	14.23	1.00	0	-17.06
45	NGC 4640	dS0,N	12 42 57.69	12 17 12.6	1931	16.80	-1.17	-0.98	14.37	1.00	0	-16.92
46	IC 810	S0	12 42 09.06	12 35 48.6	-169	15.78	-0.93	-0.65	14.41	1.00	0	-16.88
47	IC 809	dE,N	12 42 08.64	11 45 15.7	206	16.16	-1.10	-0.78	14.50	1.00	0	-16.78
48	IC 3727	Scd	12 45 05.68	10 54 03.7	85	17.43	-1.32	-1.05	14.56	1.00	0	-16.70
49	IC 783	S0/a	12 21 38.81	15 44 42.5	1293	16.96	-1.20	-1.08	14.60	1.00	0	-16.65
50	IC 3652	dE,N	12 40 58.58	11 11 04.5	470	16.44	-1.24	-0.87	14.68	1.00	0	-16.60
51	IC 3457	dE,N	12 31 51.36	12 39 25.6	1263	17.57	-1.13	-1.07	14.69	1.00	0	-16.57
52	CGCG 098-132	E/S0	12 17 27.27	17 39 02.0	894	15.44	-0.95	-0.50	14.80	1.00	0	-16.50
53	IC 3653	E	12 41 15.73	11 23 14.5	603	15.44	-0.89	-0.55	14.80	1.00	0	-16.48
54	NGC 4323	S0	12 23 01.74	15 54 20.1	1803	16.97	-1.22	-0.96	14.81	1.00	0	-16.45
55	IC 3459	dE,N	12 31 55.99	12 10 26.9	278	17.76	-1.35	-1.16	14.83	1.00	0	-16.44
56	IC 3510	dE,N	12 34 14.85	11 04 17.8	1357	16.58	-1.19	-0.80	14.87	1.00	0	-16.41
57	IC 3540	S0	12 35 27.21	12 45 00.8	753	15.89	-1.36	-0.80	14.94	1.00	0	-16.41
58	IC 3720	dE	12 44 47.44	12 03 53.4		18.53	-1.43	-1.30	14.98	1.00	1	-16.29
59	UGC 7346	dE	12 18 41.78	17 43 07.5	819	17.25	-1.40	-1.10	15.03	1.00	0	-16.27
60	UGC 7399A	dE,N	12 20 48.83	17 29 14.0	1474	16.60	-1.13	-0.91	15.04	1.00	0	-16.24
61**	VCC 723	dS0	12 24 22.08	13 01 36.9	125	12.93	-1.32	-0.39	15.14	1.00	0	-16.19
62	UGC 7436	dE	12 22 19.57	14 45 39.5	923	16.55	-1.26	-0.88	15.11	1.00	0	-16.18
63	VCC 459	BCD	12 21 11.32	17 38 19.3	2107	15.92	-1.09	-0.57	15.18	1.00	0	-16.12
64	VCC 1627	E	12 35 37.25	12 22 54.9	236	15.91	-0.89	-0.36	15.41	1.00	0	-15.91
65	IC 3486	dE,N	12 33 14.04	12 51 27.8	1903	16.76	-1.22	at edge	15.41	1.00	0	-15.87
66	UGC 7366	dE,N	12 19 28.66	17 13 50.1	925	16.59	-1.06	-0.78	15.40	1.00	0	-15.86
67	IC 3578	Scd	12 36 39.44	11 06 06.4	666	17.01	-1.14	-0.93	15.42	1.00	0	-15.86
68	IC 3292	dS0	12 24 48.38	18 11 42.6	710	16.39	-1.05	-0.68	15.43	1.00	0	-15.86
69	IC 3298	S	12 25 03.75	17 00 58.9	2452	16.65	-1.26	-0.79	15.43	1.00	0	-15.84
70	IC 3313	dE	12 25 36.45	15 49 47.6	1168	16.99	-1.30	-0.91	15.45	1.00	0	-15.83

Table 2 – continued

(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)
ID	Name	Type	α (J2000)	δ (J2000)	V_h (km s $^{-1}$)	$B(6)$	ICP	OCP	B_T	P	Class	M_B
71	IC 3665	dI	12 41 46.72	11 29 18.9	1227	17.62	-1.38	-1.08	15.44	1.00	0	-15.82
72	IC 3586	dS0	12 36 54.89	12 31 13.0	1547	16.54	-1.03	-0.68	15.58	1.00	0	-15.76
73	LSBC F644-04	dE,N	12 25 36.21	15 50 51.4		18.21	-1.45	-1.17	15.56	0.84	1	-15.72
74	IC 3461	dE,N	12 32 02.76	11 53 24.6	1038	16.70	-1.09	-0.74	15.61	1.00	0	-15.67
75	VCC 1886	dE,N	12 41 39.44	12 14 51.2	1159	17.58	-1.37	-1.02	15.65	1.00	0	-15.64
76	VCC 328	dI	12 19 11.14	12 53 09.7	2179	18.42	-1.50	-1.17	15.77	1.00	0	-15.56
77	IC 3388	dE,N	12 28 28.11	12 49 25.8	1704	17.12	-1.25	-0.86	15.72	1.00	0	-15.52
78	IC 3779	dE,N	12 47 20.65	12 09 59.5	1193	16.99	-1.17	-0.81	15.74	1.00	0	-15.51
79	VCC 1389	dE,N	12 31 52.04	12 28 54.9	936	17.51	-1.07	-0.76	15.78	1.00	0	-15.48
80	IC 783 A	S0	12 22 19.64	15 44 01.0	1159	17.22	-1.20	-0.85	15.85	1.00	0	-15.42
81	VCC 1426	dI	12 32 23.55	11 53 36.4	1110	18.46	-1.44	-1.16	15.96	1.00	0	-15.33
82	IC 3509	E	12 34 11.56	12 02 56.5	2050	17.04	-0.92	-0.62	16.01	1.00	0	-15.32
83	IC 3443	dE	12 31 15.77	12 19 54.9	1679	16.97	-0.95	-0.70	15.97	1.00	0	-15.28
84	UGC 7425	Scd	12 21 53.71	15 38 45.4	804	17.61	-1.22	-0.94	15.97	1.00	0	-15.28
85	UGC 7504	dI	12 25 21.63	16 25 47.0	913	17.29	-1.22	-0.83	15.98	1.00	0	-15.27
86*	VCC 841	BCD	12 25 47.54	14 57 08.4	503	17.04	-1.10	-0.70	16.02	1.00	0	-15.25
87	VCC 1991	dE,N	12 44 09.40	11 10 35.8		18.04	-1.14	-1.03	16.09	0.62	1	-15.21
88	VCC 530	dI	12 22 07.57	15 47 56.8	1299	18.70	-1.40	-1.17	16.07	1.00	0	-15.19
89	IC 3466	dI	12 32 05.71	11 49 04.2	903	17.07	-1.11	-0.69	16.10	1.00	0	-15.19
90	VCC 1148	E	12 28 58.13	12 39 40.2	1443	16.76	-0.68	at edge	16.06	1.00	0	-15.18
91	VCC 1561	dE,N	12 34 25.03	12 54 15.6		19.41	-1.32	-1.26	16.16	0.95	1	-15.13
92	VCC 753	dE,N	12 24 51.63	13 06 40.4		19.00	-1.34	-1.20	16.21	0.89	1	-15.09
93	VCC 1185	dE	12 29 23.55	12 27 03.4	500	17.71	-1.13	-0.91	16.18	1.00	0	-15.07
94	IC 3490	dE	12 33 13.94	10 55 43.0	80	17.81	-1.32	-0.91	16.27	1.00	0	-15.04
95	VCC 1647	dE	12 35 56.65	10 56 10.8		18.18	-1.37	-1.00	16.32	0.78	1	-14.94
96	VCC 1982	dE	12 43 51.07	11 28 01.3	938	17.57	-1.25	-0.80	16.34	1.00	0	-14.92
97	VCC 1512	S0	12 33 34.67	11 15 43.2	762	17.19	-0.91	-0.57	16.45	1.00	0	-14.92
98	VCC 797	dE,N	12 25 24.09	18 08 23.6	773	17.64	-1.19	-0.82	16.37	1.00	0	-14.91
99	VCC 1539	dE,N	12 34 06.75	12 44 30.0	1390	18.00	-1.02	-0.93	16.40	1.00	0	-14.89
100	VCC 1921	dS0	12 42 26.48	11 44 25.4		17.87	-1.24	-0.88	16.40	0.62	2	-14.88
101	VCC 1711	dE,N	12 37 22.17	12 17 14.0		17.97	-1.23	-0.90	16.46	0.71	2	-14.87
102	VCC 684	dE,N	12 23 57.74	12 53 14.0		17.92	-1.17	-0.88	16.47	0.62	2	-14.86
103	VCC 1173	dE	12 29 14.90	12 58 42.4	2468	17.60	-1.23	-0.79	16.39	1.00	0	-14.85
104	VCC 1942	dE,N	12 42 50.75	12 18 30.6		18.47	-1.29	-1.03	16.50	0.77	2	-14.79
105	UGC 7906	dI	12 44 09.85	12 06 43.5	1010	18.58	-1.24	-1.06	16.51	1.00	0	-14.75
106*	VCC 815	dE,N	12 25 37.16	13 08 37.4	-700	17.89	-1.26	-0.84	16.55	1.00	0	-14.73
107**	VCC 1995	dE	12 44 16.79	12 01 38.7		18.15	-1.36	-1.03	16.56	0.71	1	-14.69
108	VCC 846	dE,N	12 25 50.57	13 11 52.0	-730	18.05	-1.30	-0.88	16.59	1.00	0	-14.68
109	VCC 618	dI	12 23 07.44	13 44 40.3	1890	18.59	-1.34	-1.01	16.69	1.00	0	-14.67
110		VLSB	12 25 39.63	16 16 58.3		21.35	-1.28	-1.39	16.60	1.00	1	-14.66
111	VCC 1717	dE	12 37 28.98	12 21 08.9		19.27	-1.42	-1.17	16.67	0.91	1	-14.66
112	UGC 7953	dE,N	12 47 16.29	11 45 38.9		18.08	-1.35	-0.88	16.61	0.67	2	-14.65
113*	VCC 1870	dE	12 41 15.35	11 17 54.8		17.80	-1.24	-0.75	16.74	0.45	2	-14.54
114	LSBC F573-10	dE	12 22 23.79	17 01 10.9		18.93	-1.50	at edge	16.73		1	-14.53
115	VCC 1909	dE,N	12 42 07.42	11 49 42.0		18.00	-1.20	-0.81	16.76	0.56	2	-14.51
116	IC 3465	dE,N	12 32 12.29	12 03 41.8	1022	17.89	-1.14	-0.74	16.77	1.00	0	-14.51
117	VCC 1971	dE	12 43 30.93	11 02 49.8		18.02	-1.27	-0.80	16.78	0.56	2	-14.51
118	VCC 1563	dE,N	12 34 26.09	11 55 01.2		19.37	-1.22	-1.10	16.82	0.89	1	-14.51
119	VCC 539	dE,N	12 22 14.82	14 08 31.5		18.77	-1.31	-1.02	16.84	0.87	2	-14.51
120	VCC 810	dE,N	12 25 33.58	13 13 38.3	-340	18.34	-1.24	-0.90	16.78	1.00	0	-14.50
121	VCC 1149	VLSB	12 28 58.67	12 54 28.0		20.75	-1.35	-1.33	16.77	1.00	1	-14.47
122	VCC 2078	dE	12 48 43.77	11 58 11.0		19.45	-1.19	-1.17	16.84	0.93	1	-14.45
123	VCC 878	dE	12 26 10.13	14 55 44.9		19.36	-1.43	-1.14	16.87	0.87	1	-14.42
124	VCC 293	dE,N	12 18 31.89	13 11 28.4		18.84	-1.24	-1.02	16.91	0.89	1	-14.40
125	VCC 1625	VLSB	12 35 34.87	11 37 13.8		21.23	-1.21	-1.36	16.92	1.00	1	-14.40
126	VCC 1213	dE,N	12 29 39.28	12 32 54.1		18.40	-0.97	-0.91	16.88	0.79	2	-14.36
127	VCC 330	dE,N	12 19 12.47	12 51 07.2		18.66	-1.19	-0.96	16.97	0.88	2	-14.35
128	VCC 812	dE,N	12 25 34.86	15 11 40.0		18.52	-1.12	-0.93	16.96	0.85	2	-14.30
129	VCC 2081	dE	12 49 46.25	11 13 32.4		19.07	-1.41	-1.05	17.02	0.88	1	-14.29
130	VCC 515	dE	12 21 56.98	17 53 32.6		19.73	-1.42	-1.18	17.02	0.92	1	-14.28
131**	VCC 583	dE	12 22 45.13	15 30 03.0	-72	17.15	-2.15	-0.62	17.01	1.00	0	-14.26
132	VCC 1331	dE/I	12 30 58.60	11 42 28.2		20.28	-1.11	-1.26	17.05	0.92	1	-14.23
133	VCC 1915	dE	12 42 13.44	12 32 46.1		19.05	-1.49	-1.04	17.07	0.86	2	-14.22
134	VCC 1264	dE,N	12 30 10.94	12 11 44.0		18.92	-1.12	-1.02	17.04	0.89	1	-14.21
135	VCC 1663	dE	12 36 27.12	11 53 20.5		20.34	-1.45	-1.24	17.15	0.92	1	-14.21
136	VCC 769	dE	12 25 04.24	15 42 40.8		18.41	-1.32	-0.84	17.08	0.61	2	-14.18
137	VCC 1815	dE	12 39 56.42	11 54 15.7		20.17	-1.35	-1.23	17.16	0.95	1	-14.13
138*	VCC 472	dE	12 21 24.10	15 37 07.4		19.26	-1.40	-1.07	17.36	0.93	1	-14.12
139	VCC 1399	dE,N	12 32 00.80	12 37 13.2		18.34	-1.25	-0.80	17.15	0.51	2	-14.11
140**	VCC 725	dE	12 24 24.53	15 04 33.6		17.31	-1.21	-0.46	17.15	0.02	1	-14.10

Table 2 – continued

(1) ID	(2) Name	(3) Type	(4) α (J2000)	(5) δ (J2000)	(6) V_h (km s $^{-1}$)	(7) $B(6)$	(8) ICP	(9) OCP	(10) B_T	(11) P	(12) Class	(13) M_B
141	VCC 554	dE,N	12 22 24.34	15 28 15.8		19.33	-1.29	-1.08	17.17	0.92	1	-14.10
142	VCC 1879	dE,N	12 41 27.38	11 08 45.5		19.20	-1.22	-1.04	17.19	0.92	1	-14.09
143	VCC 1313	BCD	12 30 48.52	12 02 42.1	1254	17.27	-0.71	-0.11	17.18	1.00	0	-14.09
144	VCC 793	dI	12 25 21.31	13 04 14.4	1908	18.33	-1.34	-0.89	17.20	1.00	0	-14.09
145	VCC 1396	dE,N	12 31 56.43	11 58 22.2		19.52	-1.17	-1.12	17.21	0.88	1	-14.06
146	VCC 594	dE	12 22 51.14	15 16 30.7		19.36	-1.40	-1.08	17.23	0.92	1	-14.04
147	VCC 1565	dE,N	12 34 30.47	11 44 04.4		19.20	-1.10	-1.09	17.29	0.93	1	-14.03
148	VCC 1418	dE,N	12 32 11.39	12 30 25.4		19.02	-1.25	-0.98	17.25	0.90	2	-14.02
149	VCC 2011	dE	12 45 04.23	12 21 03.8		19.26	-1.48	-1.03	17.27	0.92	1	-14.01
150	VCC 1123	dE,N	12 28 42.67	12 32 59.9		18.96	-1.05	-0.96	17.26	0.90	1	-13.99
151	VCC 663	dE	12 23 42.92	18 39 43.1		18.90	-1.29	-0.95	17.30	0.87	2	-13.99
152		dI	12 51 06.88	12 03 39.4		18.83	-1.40	-0.92	17.31	0.87	2	-13.99
153	VCC 350	dI	12 19 26.00	13 18 38.4	305	19.15	-1.32	-0.99	17.34	1.00	0	-13.99
154	VCC 1599	dE	12 35 06.60	11 54 03.1		20.11	-1.46	-1.19	17.36	0.93	1	-13.99
155	VCC 1891	dE,N	12 41 48.94	11 11 29.5		18.32	-1.14	-0.73	17.30	0.43	2	-13.97
156	VCC 668	dE	12 23 47.10	15 07 32.0		19.26	-1.38	-1.03	17.30	0.92	1	-13.95
157	VCC 1369	dE,N	12 31 33.39	12 03 49.8		18.50	-1.08	-0.76	17.33	0.43	2	-13.94
158	VCC 1551	VLSB	12 34 15.38	11 28 01.7		21.35	-1.44	-1.33	17.38	1.00	1	-13.93
159		dI	12 25 46.29	16 38 07.8		18.54	-1.10	-0.79	17.33	0.51	2	-13.92
160	VCC 748	dE	12 24 47.59	14 34 35.6		19.11	-1.36	-0.96	17.39	0.91	2	-13.90
161		VLSB	12 32 34.62	12 38 15.4		21.10	-1.67	-1.45	17.38	1.00	1	-13.89
162	VCC 872	dE,N	12 26 06.72	12 51 40.0	1265	18.47	-0.97	-0.73	17.41	1.00	0	-13.87
163**	VCC 802	BCD	12 25 28.72	13 29 51.5	-215	18.34	-0.90	-0.60	17.45	1.00	0	-13.87
164*	VCC 422	dE	12 20 30.13	18 19 16.0		19.68	-1.43	-1.10	17.46	0.91	1	-13.85
165	VCC 1352	dE	12 31 19.60	12 36 41.9		18.63	-1.19	-0.80	17.42	0.54	2	-13.83
166	VCC 1858	dE	12 40 54.17	12 31 56.3		20.29	-1.32	-1.20	17.48	0.92	1	-13.83
167	VCC 1951	dE,N	12 43 02.26	11 41 53.3		19.03	-1.03	-0.93	17.45	0.91	1	-13.82
168	VCC 2088	dE,N	12 51 11.12	11 14 39.0		19.85	-1.05	-1.14	17.49	0.91	1	-13.82
169	VCC 761	dE	12 25 00.28	15 36 15.8		19.36	-1.30	-1.02	17.44	0.93	1	-13.81
170	VCC 1606	dE,N	12 35 14.74	12 14 15.3		18.90	-0.90	-0.86	17.51	0.76	2	-13.80
171		dE	12 31 36.86	11 00 28.9		19.29	-1.33	-0.99	17.50	0.93	1	-13.79
172	VCC 1366	dE,N	12 31 31.72	11 36 11.3		19.31	-0.96	-0.98	17.56	0.93	1	-13.78
173	VCC 2032	dE	12 45 52.39	11 14 45.9		19.91	-1.25	-1.13	17.52	0.95	1	-13.77
174	VCC 833	dE,N	12 25 44.66	13 01 20.0	720	18.81	-1.00	-0.82	17.53	1.00	0	-13.75
175	VCC 1798	VLSB	12 39 31.27	11 27 14.3		21.57	-1.42	-1.34	17.56	1.00	1	-13.74
176	VCC 646	dE	12 23 31.80	17 47 40.5		19.09	-1.27	-0.91	17.55	0.89	2	-13.73
177		VLSB	12 44 22.28	12 00 34.1		21.35	-1.03	-1.34	17.55	1.00	1	-13.70
178**	VCC 1689	dE,N	12 36 51.32	12 22 08.7		19.49	-1.43	-1.34	17.65	1.00	1	-13.69
179	VCC 2062	dI	12 48 00.03	10 58 14.7	1140	20.07	-1.40	-0.91	17.64	1.00	0	-13.68
180		VLSB	12 20 51.26	16 21 49.6		21.32	-1.47	-1.31	17.60	1.00	1	-13.67
181	VCC 1905	dE	12 42 03.13	12 28 50.8		20.52	-1.42	-1.21	17.65	0.90	1	-13.65
182	VCC 1464	dE	12 32 53.87	11 11 28.8		19.51	-1.40	-0.99	17.72	0.95	1	-13.64
183**	VCC 1403	dE/I	12 32 00.37	13 04 58.3		19.80	-1.57	-3.84	17.64	1.00	1	-13.63
184	VCC 1681	dE,N	12 36 37.44	11 09 13.1		19.72	-1.13	-1.05	17.67	0.93	1	-13.62
185	VCC 863	dE	12 25 59.75	14 02 22.5		20.37	-1.24	-1.17	17.74	0.93	1	-13.62
186**	VCC 1683	dE	12 36 38.46	10 56 15.9		15.33	-2.61	-0.42	17.65	0.46	1	-13.60
187		VLSB	12 23 22.16	19 12 11.5		22.01	-1.51	-1.36	17.75	1.00	1	-13.58
188	VCC 677	dE	12 23 53.34	18 37 56.8		19.04	-1.33	-0.85	17.72	0.71	2	-13.56
189	VCC 696	dE	12 24 04.04	17 32 57.5		19.15	-1.36	-0.89	17.72	0.79	2	-13.55
190	VCC 1191	dE,N	12 29 28.72	12 29 47.0		19.11	-1.11	-0.87	17.71	0.74	2	-13.54
191	VCC 795	dE,N	12 25 23.18	14 48 12.9		18.81	-1.18	-0.72	17.77	0.57	2	-13.51
192*	VCC 625	dE	12 23 11.25	14 51 45.0		19.17	-1.26	-0.85	17.79	0.69	2	-13.50
193	VCC 818	dI	12 25 37.79	16 39 51.8		19.59	-1.38	-0.98	17.80	0.95	1	-13.45
194	VCC 813	dE,N	12 25 35.75	16 35 46.1		19.81	-0.98	-1.04	17.80	0.95	1	-13.45
195*	VCC 1517	dE,N	12 33 40.85	12 34 17.0		19.49	-1.11	-0.95	17.82	0.97	1	-13.45
196**	VCC 716	dE	12 24 13.70	14 55 44.6		18.97	-1.31	-1.73	17.84	1.00	1	-13.42
197	VCC 877	dE,N	12 26 09.58	13 40 23.7		19.04	-1.23	-0.77	17.89	0.64	2	-13.42
198*	VCC 779	dE,N	12 25 13.15	13 01 32.0		19.55	-0.97	-0.95	18.09	0.98	1	-13.42
199		dE	12 23 05.16	15 55 54.5		19.60	-1.38	-0.98	17.85	0.95	1	-13.41
200**	VCC 454	dE	12 21 05.45	15 43 13.4		19.71	-1.44	-1.18	17.84	0.93	1	-13.41
201	VCC 1609	dE,N	12 35 20.34	11 38 10.6		19.16	-0.92	-0.77	17.91	0.63	1	-13.41
202	VCC 861	dE	12 25 58.93	15 16 37.7		19.08	-1.31	-0.80	17.85	0.63	2	-13.40
203	VCC 1642	dE,N	12 35 53.17	11 40 55.4		19.52	-0.78	-0.93	17.91	0.97	1	-13.40
204	VCC 1413	dE	12 32 07.71	12 26 03.0		19.80	-1.29	-1.02	17.89	0.95	1	-13.38
205	VCC 1153	dE	12 28 59.85	12 38 55.0		19.09	-1.28	-0.80	17.88	0.63	2	-13.36
206	VCC 1216	dE	12 29 41.36	12 02 47.5		20.41	-1.26	-1.14	17.92	0.89	1	-13.35
207		dE	12 17 32.66	18 24 17.3		19.76	-1.52	-0.99	17.96	0.95	1	-13.32
208	VCC 650	dE,N	12 23 34.10	13 17 59.1		19.32	-0.71	-0.79	18.14	0.78	1	-13.30
209	VCC 674	dE,N	12 23 52.68	13 52 56.9		19.25	-0.74	-0.79	18.04	0.71	1	-13.28
210	VCC 2001	dE	12 44 33.77	11 47 42.9		19.67	-1.34	-0.95	18.00	0.98	1	-13.26

Table 2 – continued

(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)
ID	Name	Type	α (J2000)	δ (J2000)	V_h (km s $^{-1}$)	$B(6)$	ICP	OCP	B_T	P	Class	M_B
211	VCC 777	dE	12 25 11.34	14 26 29.1		19.82	-1.33	-0.98	18.05	0.97	1	-13.26
212	VCC 1259	dE	12 30 06.14	12 22 38.1		19.70	-1.40	-0.96	18.01	0.97	1	-13.25
213	VCC 1794	dE,N	12 39 27.07	11 46 35.3		18.99	-1.15	-0.64	18.08	0.66	2	-13.24
214	VCC 1136	dE,N	12 28 49.05	12 07 57.7		20.31	-1.09	-1.10	18.04	0.94	1	-13.23
215		dI	12 24 45.54	13 20 25.6		19.42	-1.22	-0.82	18.11	0.77	2	-13.22
216	VCC 1783	dE	12 39 13.50	12 06 03.0		19.78	-1.07	-0.94	18.14	0.97	1	-13.22
217	VCC 726	dE,N	12 24 24.14	16 23 13.8		20.42	-0.96	-1.12	18.06	0.93	1	-13.20
218*	VCC 378	dE	12 19 50.57	15 40 17.1		20.24	-1.41	-1.09	18.05	0.97	1	-13.20
219	VCC 1131	dE	12 28 45.81	12 01 18.8		19.97	-1.44	-1.01	18.07	0.98	1	-13.20
220	VCC 1785	dE,N	12 39 15.73	11 16 11.7		19.54	-1.01	-0.88	18.09	0.84	1	-13.20
221	VCC 1831	dE,N	12 40 18.34	10 59 47.7		19.93	-1.17	-1.00	18.09	0.98	1	-13.16
222	VCC 2003	dE	12 44 41.47	11 31 13.6		19.90	-1.21	-0.96	18.11	1.00	1	-13.16
223		VLSB	12 36 35.18	11 37 03.9		21.32	-1.02	-1.25	18.15	1.00	1	-13.16
224	VCC 1454	dE,N	12 32 44.95	10 56 58.5		20.31	-1.08	-1.07	18.18	0.97	1	-13.15
225		dI	12 41 12.07	10 55 59.3		19.15	-1.12	-0.69	18.17	0.61	2	-13.08
226	VCC 780	dE	12 25 13.24	14 50 44.7		19.77	-1.30	-0.93	18.19	0.97	1	-13.08
227*		VLSB	12 25 22.24	19 40 28.9		21.38	-1.13	-1.21	18.22	0.90	2	-13.08
228		dE	12 22 59.49	16 58 58.7		18.95	-1.15	-0.58	18.20	0.45	2	-13.06
229		dE	12 24 36.57	18 56 15.2		20.78	-1.30	-1.12	18.30	0.93	1	-13.00
230		dE	12 19 23.32	19 35 14.2		19.34	-1.28	-0.73	18.28	0.75	2	-12.98
231	VCC 1736	dE	12 37 55.18	11 08 56.3		20.78	-1.40	-1.14	18.33	0.90	1	-12.94
232	VCC 1278	dE	12 30 17.50	12 14 28.3		20.36	-1.41	-1.04	18.32	0.98	1	-12.93
233	VCC 1815	dE	12 39 57.72	11 54 26.0		21.07	-1.38	at edge	18.38		1	-12.92
234**	VCC 635	dE	12 23 21.99	13 20 38.4		19.43	-1.21	-1.21	18.56	0.96	2	-12.89
235	VCC 1637	dE,N	12 35 45.56	12 10 53.3		19.44	-0.99	-0.69	18.45	0.78	2	-12.88
236	VCC 714	dE	12 24 13.49	17 30 16.1		19.41	-1.29	-0.70	18.41	0.80	2	-12.86
237*		VLSB	12 44 54.57	10 57 44.6		21.49	-1.30	-1.24	18.42	0.84	1	-12.85
238	VCC 1880	dE	12 41 28.03	12 25 40.4		19.57	-1.31	-0.75	18.45	0.78	1	-12.85
239	VCC 757	dE	12 24 53.10	14 39 19.1		19.51	-1.36	-0.71	18.47	0.77	2	-12.84
240*	VCC 1700	dE	12 37 03.30	11 28 42.9		21.14	-1.34	-1.17	18.48	0.91	1	-12.83
241**	VCC 505	dE,N	12 21 48.99	18 25 45.0		20.35	-1.41	-1.41	18.49	1.00	1	-12.83
242	VCC 1754	dE	12 38 17.12	11 10 51.1		20.67	-1.40	-1.09	18.45	0.97	1	-12.82
243	VCC 1402	dE,N	12 32 00.28	11 01 24.1		19.86	-1.25	-0.85	18.50	0.88	1	-12.82
244		dE,N	12 20 11.13	17 43 05.1		19.80	-1.33	-0.81	18.53	0.86	2	-12.78
245	VCC 521	dE	12 22 2.84	17 12 03.8		20.09	-1.40	-0.92	18.53	0.97	1	-12.73
246		dE	12 41 05.03	12 15 58.1		19.76	-1.27	-0.78	18.58	0.87	1	-12.73
247*	VCC 850	dI	12 25 52.79	13 11 32.7		19.97	-1.28	-1.00	18.56	1.00	2	-12.71
248	VCC 644	dE	12 23 28.89	17 32 25.0		19.73	-1.35	-0.75	18.58	0.86	2	-12.71
249	VCC 1595	dE	12 34 59.99	11 32 50.0		19.73	-1.26	-0.74	18.64	0.86	1	-12.71
250	VCC 2025	dE	12 45 35.37	11 33 07.5		20.02	-1.23	-0.82	18.58	0.88	1	-12.68
251	VCC 742	dE,N	12 24 40.40	15 42 35.9		20.12	-0.83	-0.91	18.58	0.93	1	-12.68
252	VCC 1613	dE	12 35 26.88	12 31 41.3		19.81	-1.27	-0.77	18.64	0.87	1	-12.68
253**	VCC 624	dE,N	12 23 11.66	13 25 06.8		20.55	-1.01	-0.99	18.75	1.00	2	-12.67
254*	VCC 519	dE	12 22 00.87	14 08 11.6		20.52	-1.37	-1.00	18.70	1.00	1	-12.66
255**	VCC 721	dI	12 24 22.02	13 25 04.6		20.01	-0.97	-0.83	18.69	0.90	2	-12.65
256		dE,N	12 35 59.74	11 27 08.3		20.39	-0.94	-0.97	18.66	1.00	2	-12.65
257*	VCC 1594	dE	12 35 00.08	11 20 38.3		19.58	-1.23	-0.67	18.65	0.81	2	-12.64
258	VCC 1986	dE,N	12 43 57.80	11 52 50.7		20.24	-0.92	-0.94	18.61	0.98	1	-12.64
259	VCC 1963	dE,N	12 43 18.07	11 28 30.6		20.88	-0.70	-1.10	18.64	0.97	1	-12.62
260	VCC 789	dE	12 25 19.16	13 15 24.0		20.36	-1.33	-0.96	18.68	1.00	1	-12.62
261	VCC 1518	dE	12 33 40.89	12 22 56.7		20.16	-1.33	-0.90	18.67	0.93	1	-12.61
262	VCC 1729	dE	12 37 46.06	10 59 07.1		19.79	-1.30	-0.77	18.65	0.86	2	-12.60
263	VCC 1522	dE,N	12 33 47.08	11 46 54.4		20.83	-0.89	-1.06	18.75	0.96	1	-12.57
264	VCC 1718	dE,N	12 37 30.13	11 28 54.4		19.83	-1.10	-0.74	18.74	0.87	1	-12.57
265	VCC 1405	dE	12 32 00.05	11 18 06.9		21.19	-1.33	-1.12	18.81	0.92	1	-12.56
266		VLSB	12 21 52.38	17 29 58.7		21.59	-1.39	-1.21	18.73	0.89	2	-12.55
267		dE	12 38 38.84	11 28 53.8		19.35	-1.13	-0.48	18.77	0.51	2	-12.51
268	VCC 708	dE,N	12 24 13.41	13 37 57.5		19.86	-1.23	-0.71	18.83	0.87	2	-12.50
269	VCC 1536	dE,N	12 34 06.59	11 50 12.4		20.38	-0.69	-0.92	18.82	0.98	1	-12.49
270*	VCC 1381	dE	12 31 44.05	12 36 45.2		20.12	-1.33	-0.87	18.78	0.92	1	-12.48
271		dI	12 30 24.03	12 26 09.5		21.31	-1.36	-1.13	18.76	0.95	1	-12.48
272**	VCC 1746	dE	12 38 11.83	12 03 29.7		20.34	-1.50	-1.93	18.86	1.00	1	-12.46
273	VCC 845	dE	12 25 48.16	13 51 15.1		20.16	-1.19	-0.83	18.86	0.90	2	-12.46
274	VCC 719	dE	12 24 18.78	12 54 45.5		20.00	-1.23	-0.75	18.90	0.90	1	-12.43
275	VCC 557	dE	12 22 29.14	13 18 56.4		19.67	-1.19	-0.54	18.98	0.56	2	-12.43
276	VCC 830	dE	12 25 42.55	17 59 25.2		20.05	-1.25	-0.59	18.84	0.72	2	-12.42
277	VCC 643	dE	12 23 29.45	14 53 22.2		20.01	-1.27	-0.76	18.86	0.89	1	-12.42
278**		dE/I	12 29 19.59	12 22 37.7		20.06	-1.20	-3.14	18.94	1.00	1	-12.41
279	VCC 2010	dE	12 44 53.04	12 10 58.9		19.99	-1.29	-0.75	18.87	0.90	1	-12.40
280	VCC 432	dE	12 20 46.73	17 13 52.0		19.63	-1.21	-0.56	18.89	0.57	2	-12.37

Table 2 – continued

(1) ID	(2) Name	(3) Type	(4) α (J2000)	(5) δ (J2000)	(6) V_h (km s $^{-1}$)	(7) $B(6)$	(8) ICP	(9) OCP	(10) B_T	(11) P	(12) Class	(13) M_B
281		dE/I	12 30 47.17	11 32 16.9		20.37	-1.41	-0.87	18.96	0.95	1	-12.37
282		VLSB	12 51 01.10	11 29 01.4		21.66	-1.33	-1.17	19.01	0.91	1	-12.36
283	VCC 1621	dE	12 35 34.08	11 47 10.1		19.92	-1.26	-0.68	18.97	0.91	2	-12.35
284	VCC 1925	dE	12 42 33.77	11 49 14.6		21.43	-1.26	-1.14	18.95	0.95	1	-12.32
285		VLSB	12 24 08.57	13 49 59.3		21.81	-1.28	-1.20	18.99	1.00	1	-12.32
286		dE/I	12 40 04.74	11 34 15.4		20.03	-1.21	-0.69	18.99	0.88	1	-12.31
287	VCC 1841	dE	12 40 27.05	12 11 57.0		20.22	-1.31	-0.80	19.05	0.92	2	-12.29
288	IBM88 V03L12	dE,N	12 24 35.15	15 09 50.7		20.60	-1.31	-0.94	18.99	1.00	1	-12.28
289		dE/I	12 19 50.59	16 16 08.1		20.08	-1.22	-0.74	19.00	0.90	1	-12.27
290	VCC 1416	dE,N	12 32 10.30	12 33 02.6		21.26	-0.68	-1.14	19.00	0.96	1	-12.26
291	VCC 2072	dE	12 48 25.23	12 14 15.7		20.52	-1.36	-0.89	19.03	0.97	1	-12.24
292		dE,N	12 45 24.97	10 55 26.6		20.51	-0.82	-0.89	19.03	0.96	1	-12.23
293		dI	12 24 12.90	14 29 39.1		20.21	-1.32	-0.76	19.08	0.92	1	-12.22
294*		dE/I	12 32 07.37	11 20 31.5		21.08	-1.20	-1.02	19.16	0.97	1	-12.22
295	VCC 600	dE	12 22 55.51	15 33 34.4		20.49	-1.14	-0.87	19.06	0.97	1	-12.20
296**	VCC 704	dE	12 24 11.32	13 22 25.1		20.70	-1.21	-1.28	19.17	0.85	1	-12.18
297	VCC 1143	dE,N	12 28 55.69	12 42 29.3		20.05	-1.31	at edge	19.07		1	-12.17
298**		dE	12 40 14.56	11 31 57.2		20.28	-1.49	-3.56	19.15	1.00	1	-12.17
299	VCC 2047	dE	12 47 08.59	11 27 39.9		19.83	-1.17	-0.55	19.12	0.67	1	-12.15
300	VCC 1161	dE	12 29 05.45	12 01 52.9		20.25	-1.29	-0.76	19.11	0.95	1	-12.15
301	VCC 1904	dE,N	12 42 02.76	10 57 13.4		19.98	-0.43	-0.64	19.11	0.85	1	-12.14
302	VCC 754	dE	12 24 50.72	15 00 38.9		20.23	-1.33	-0.74	19.11	0.92	1	-12.14
303	VCC 536	dE	12 22 12.26	16 58 28.0		20.51	-1.33	-0.86	19.12	0.97	1	-12.14
304		dE/I	12 25 55.13	19 11 51.1		20.51	-1.09	-0.86	19.16	0.97	2	-12.14
305		dE	12 20 35.05	18 53 02.8		20.34	-1.31	-0.76	19.19	0.92	1	-12.12
306*		dI	12 25 11.26	13 27 56.0		20.60	-1.32	-0.86	19.22	0.96	2	-12.11
307*	VCC 1990	dE	12 44 06.81	12 41 03.9		20.94	-1.47	-0.98	19.17	1.00	1	-12.10
308	VCC 495	dE	12 21 41.78	17 49 40.7		20.24	-1.25	-0.67	19.22	0.90	2	-12.08
309		VLSB	12 37 37.71	12 28 38.5		22.23	-1.27	-1.22	19.27	7.00	1	-12.08
310**		VLSB	12 31 48.55	10 58 09.3		21.20	-1.01	-1.96	19.22	1.00	1	-12.08
311	VCC 547	dE	12 22 20.91	15 09 35.2		20.29	-1.31	-0.74	19.21	0.90	1	-12.06
312	VCC 1271	dE	12 30 15.31	12 30 58.1		20.50	-1.30	-0.82	19.20	0.94	1	-12.05
313		dI	12 25 20.95	13 49 09.4		21.04	-1.30	-0.98	19.27	1.00	1	-12.05
314	VCC 1157	dE	12 29 02.05	12 26 05.5		20.73	-1.23	-0.90	19.21	0.96	1	-12.04
315		VLSB	12 24 26.20	13 28 05.9		21.80	-1.32	-1.11	19.26	1.00	1	-12.04
316	VCC 829	dE	12 25 42.78	15 34 30.1		20.14	-1.23	-0.63	19.22	0.82	2	-12.04
317	VCC 1461	dE,N	12 32 51.28	11 17 45.2		20.62	-1.13	-0.81	19.36	0.93	1	-12.03
318		VLSB	12 22 09.57	15 39 10.7		21.27	-1.34	-1.05	19.24	1.00	1	-12.02
319	VCC 605	dE,N	12 23 02.05	13 33 33.1		20.60	-1.11	-0.83	19.39	0.95	2	-12.00
320**	VCC 844	dE	12 25 48.40	13 07 21.5		20.33	-1.29	-1.46	19.28	1.00	1	-12.00
321		VLSB	12 46 48.33	12 11 49.5		21.77	-1.24	-1.15	19.28	1.00	1	-11.99
322		dI	12 43 08.14	11 05 55.3		21.52	-1.19	-1.09	19.29	0.85	2	-11.99
323		dE,N	12 34 01.41	12 43 11.1		20.67	-1.17	-0.85	19.33	0.95	1	-11.96
324		dE/I	12 44 54.13	11 01 08.3		20.84	-1.34	-0.90	19.34	0.96	1	-11.96
325		dI	12 23 39.13	13 49 04.8		20.20	-1.30	-0.61	19.37	0.78	2	-11.95
326	VCC 1672	dE,N	12 36 32.63	12 31 05.9		20.23	-0.44	-0.61	19.41	7.73	1	-11.94
327	VCC 804	dE,N	12 25 30.69	12 58 37.9		21.03	-1.14	-0.96	19.35	1.00	1	-11.94
328		dE,N	12 23 53.17	13 30 23.6		20.48	-0.95	-0.72	19.43	0.88	2	-11.93
329		dE,N	12 33 51.17	12 57 30.8		20.34	-1.00	-0.67	19.41	0.87	1	-11.88
330	VCC 1466	dE	12 32 55.37	12 38 06.9		20.21	-1.23	-0.61	19.40	0.73	2	-11.87
331*		dE	12 35 19.70	11 29 20.8		21.00	-1.10	-0.90	19.50	0.94	1	-11.84
332*		dI	12 21 13.27	16 17 36.9		20.27	-1.05	-0.64	19.44	0.86	2	-11.83
333		dE	12 45 16.32	12 14 53.7		21.05	-1.21	-0.84	19.45	0.91	1	-11.82
334	VCC 647	dE	12 23 33.43	17 49 14.9		20.73	-1.38	-0.82	19.46	0.94	1	-11.82
335	VCC 1558	dI	12 34 21.35	11 24 57.7		20.42	-1.24	-0.65	19.51	0.88	2	-11.80
336	VCC 1538	dE	12 34 05.98	11 03 15.7		20.74	-1.38	-0.81	19.49	0.94	1	-11.79
337		dE/I	12 24 34.34	13 22 24.4		20.99	-1.43	-0.88	19.56	0.94	1	-11.79
338	VCC 603	dE	12 22 59.28	13 45 25.0		20.40	-1.29	-0.62	19.57	0.80	2	-11.79
339		VLSB	12 45 34.62	11 28 49.2		21.34	-1.32	-1.00	19.49	1.00	1	-11.77
340		dE/I	12 48 00.30	11 22 21.5		20.95	-1.24	-0.88	19.51	0.96	1	-11.77
341		dE,N	12 44 45.02	11 48 04.1		20.82	-0.80	-0.84	19.50	0.92	2	-11.76
342	IBM88 V03L15	dE/I	12 25 55.00	14 38 28.5		20.71	-1.26	-0.77	19.54	0.94	1	-11.75
343		VLSB	12 43 17.11	11 37 41.4		21.46	-0.90	-1.02	19.53	1.00	1	-11.74
344	IBM88 V07L04	dI	12 29 53.69	12 37 13.1		21.45	-1.15	-1.16	19.51	1.00	1	-11.73
345**	VCC 1103	dE,N	12 28 26.29	12 20 45.6		21.17	-0.71	-1.00	19.57	0.99	1	-11.70
346*		dE/I	12 43 08.00	11 37 16.6		21.10	-1.41	-0.91	19.58	0.98	1	-11.69
347		dI	12 17 49.17	16 35 46.8		20.60	-1.28	-0.68	19.58	0.89	2	-11.68
348		dE/I	12 26 08.19	15 28 53.6		20.37	-1.22	-0.57	19.61	0.73	2	-11.65
349	VCC 1277	dE	12 30 18.02	12 02 30.4		20.83	-1.35	-0.79	19.63	0.94	1	-11.65
350	VCC 1968	dE	12 43 27.26	10 57 23.0		20.34	-1.32	-0.55	19.63	0.70	1	-11.63

Table 2 – continued

(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)
ID	Name	Type	α (J2000)	δ (J2000)	V_h (km s $^{-1}$)	$B(6)$	ICP	OCP	B_T	P	Class	M_B
351		dE/I	12 50 35.95	11 23 18.6		20.44	-1.41	-0.61	19.71	0.81	2	-11.62
352	VCC 1882	dE,N	12 41 30.80	11 40 56.2		20.81	-0.60	-0.76	19.68	0.90	2	-11.61
353	VCC 556	dE	12 22 28.10	13 09 46.1		20.70	-1.20	-0.68	19.75	0.90	2	-11.61
354		dI	12 30 18.23	12 34 18.3		20.96	-1.31	-0.83	19.64	0.94	1	-11.60
355		dE/I	12 31 36.45	13 05 19.4		20.99	-1.31	at edge	19.67		1	-11.59
356	VCC 1500	dE	12 33 22.56	11 38 29.9		20.40	-1.29	-0.62	19.78	0.79	2	-11.59
357		dE	12 44 52.86	11 33 34.1		20.69	-1.30	-0.71	19.68	0.88	1	-11.58
358		VLSB	12 36 18.14	11 57 12.1		21.71	-1.48	-1.01	19.82	1.00	1	-11.56
359		dE/I	12 50 45.45	11 54 45.5		20.84	-1.24	-0.72	19.77	0.86	1	-11.55
360		VLSB	12 23 54.55	13 10 56.2		21.64	-1.14	-0.99	19.85	1.00	1	-11.54
361		dE,N	12 42 22.81	11 41 07.2		21.34	-1.09	-0.90	19.76	0.98	1	-11.53
362	VCC 767	dE	12 25 04.80	13 04 32.3		21.21	-1.26	-0.88	19.76	1.00	1	-11.53
363	VCC 1977	dE	12 43 38.41	11 17 51.9		20.52	-1.26	-0.58	19.75	0.77	2	-11.52
364	VCC 1769	dE	12 38 38.03	12 36 39.2		20.73	-1.35	-0.67	19.80	0.90	1	-11.52
365	VCC 1083	dE	12 28 12.24	11 58 13.4		20.33	-1.22	-0.46	19.76	0.64	2	-11.51
366		dI	12 36 13.25	12 10 10.9		21.61	-1.22	-0.98	19.85	1.00	2	-11.50
367	VCC 1162	dE,N	12 29 05.19	12 09 14.0		20.93	-0.71	-0.77	19.78	0.92	2	-11.49
368**		VLSB	12 23 11.67	15 23 9.4		21.40	-1.33	-1.35	19.79	1.00	1	-11.48
369		dE/I	12 33 07.56	12 12 13.7		21.23	-1.02	-0.87	19.81	0.95	2	-11.48
370	VCC 1494	dE	12 33 16.93	12 16 57.2		20.93	-1.43	-0.73	19.82	0.88	1	-11.46
371		dI	12 42 48.07	12 38 48.0		21.19	-1.11	-0.85	19.82	0.95	2	-11.46
372		dE/I	12 22 19.99	15 40 46.8		20.53	-1.08	-0.55	19.82	0.72	2	-11.45
373	VCC 1578	dE	12 34 41.76	11 08 34.2		20.67	-1.30	-0.56	19.88	0.77	2	-11.38
374		VLSB	12 18 50.72	15 54 18.4		21.79	-1.37	-1.01	19.91	1.00	1	-11.36
375	VCC 1680	dE	12 36 36.69	10 59 28.6		20.47	-1.18	-0.46	19.90	0.66	2	-11.35
376	VCC 607	dE	12 23 02.28	13 54 50.1		20.91	-1.35	-0.64	20.00	0.87	1	-11.34
377		dI	12 23 49.02	15 14 39.8		20.94	-1.19	-0.70	19.94	0.85	2	-11.33
378**	VCC 1286	dE	12 30 24.61	12 47 35.2		20.57	-1.41	-0.94	19.92	0.99	1	-11.32
379		dE	12 25 37.54	17 50 36.9		20.62	-1.15	-0.53	19.94	0.77	2	-11.32
380		dI	12 46 17.91	11 11 09.4		20.91	-1.51	-0.66	19.99	0.90	1	-11.32
381		dI	12 35 07.06	11 39 37.6		21.61	-1.24	-0.93	20.00	1.00	2	-11.32
382*		dE/I	12 22 39.27	18 05 20.1		20.98	-1.45	-0.69	20.01	0.90	2	-11.29
383		dE,N	12 18 14.47	16 44 08.2		21.57	-1.08	-0.92	19.99	0.95	1	-11.27
384	VCC 1634	dE	12 35 41.86	12 12 24.4		21.46	-0.94	-0.84	20.05	1.00	1	-11.27
385		dE/I	12 32 19.36	10 56 29.7		21.65	-1.11	-0.93	20.06	1.00	1	-11.27
386		VLSB	12 32 30.07	11 50 01.4		21.65	-1.29	-0.94	20.03	0.94	1	-11.26
387**		dI	12 44 35.23	12 18 34.4		21.68	-1.24	-1.44	20.68	1.00	2	-11.26
388		dI	12 30 57.44	11 05 15.9		20.77	-1.31	-0.55	20.06	0.78	2	-11.23
389		dE/I	12 24 51.16	15 23 39.4		20.93	-1.31	-0.71	20.04	0.88	1	-11.22
390		dE	12 44 39.14	12 19 14.7		20.66	-1.12	-0.44	20.06	0.76	2	-11.21
391		dE,N	12 37 15.99	11 39 47.4		20.89	-1.01	-0.59	20.11	0.82	2	-11.20
392		dI	12 21 54.17	13 15 01.8		20.66	-1.08	-0.40	20.19	0.81	2	-11.20
393**		dI	12 34 14.76	12 42 55.2		21.45	-1.45	-1.64	20.09	1.00	1	-11.20
394**		dE/I	12 40 01.63	11 52 45.3		20.97	-0.71	-0.84	20.09	0.93	1	-11.20
395**		dE,N	12 28 46.97	12 38 31.7		21.39	-0.82	-1.05	20.05	0.98	2	-11.19
396	VCC 1635	dE	12 35 40.82	12 14 07.6		20.89	-1.17	-0.58	20.13	0.84	2	-11.19
397		VLSB	12 22 22.52	14 25 50.3		21.81	-1.37	-0.94	20.18	1.00	1	-11.13
398**		dI	12 30 28.27	12 58 57.6		20.80	-1.24	-1.91	20.14	1.00	2	-11.10
399		dE/I	12 49 04.28	11 10 37.2		21.66	-1.18	-0.88	20.21	0.92	2	-11.07
400**	VCC 1139	dE	12 28 51.25	11 57 28.1		20.72	-1.26	-1.39	20.21	1.00	1	-11.05
401		VLSB	12 20 24.62	16 04 22.2		21.26	-0.64	-0.71	20.24	0.80	1	-11.03
402		dI	12 41 14.05	12 14 58.8		21.33	-1.22	-0.72	20.28	0.81	1	-11.02
403**		VLSB	12 42 53.94	12 32 16.5		21.70	-1.21	-1.03	20.28	1.00	1	-11.00
404		dI	12 40 37.18	11 07 26.6		20.77	-1.16	-0.42	20.26	0.83	2	-11.00
405**		dE/I	12 44 24.22	12 10 25.4		21.48	-0.72	-0.62	20.27	0.83	2	-10.99
406	VCC 1631	dE	12 35 38.34	12 20 29.9		20.88	-1.26	-0.42	20.33	0.85	1	-10.99
407		dE	12 39 57.47	12 06 47.4		21.13	-1.28	-0.57	20.38	0.82	1	-10.98
408**		dI	12 37 19.38	11 52 12.5		21.22	-1.24	-1.71	20.37	1.00	1	-10.98
409		dE	12 33 45.32	10 52 19.7		21.17	-1.18	-0.63	20.31	0.88	2	-10.96
410		dE/I	12 47 27.97	12 11 50.7		20.99	-1.25	-0.52	20.32	0.84	2	-10.94
411		dE/I	12 30 53.75	10 54 43.0		21.48	-1.25	-0.76	20.36	0.87	1	-10.93
412**		dE/I	12 37 03.28	11 25 09.9		21.66	-1.54	-3.34	20.37	1.00	1	-10.93
413**		dI	12 31 28.20	12 51 21.4		21.10	-1.35	-1.49	20.33	1.00	1	-10.92
414*		dE	12 24 27.85	18 27 24.5		21.39	-1.17	-0.69	20.41	0.85	2	-10.90
415*		dE/I	12 45 10.07	11 41 20.2		21.04	-1.02	-0.56	20.36	0.80	2	-10.90
416		dE/I	12 23 06.05	17 05 25.1		21.33	-1.14	-0.67	20.40	0.85	2	-10.86
417		dI	12 40 22.80	11 13 36.9		21.50	-1.05	-0.74	20.42	0.86	2	-10.86
418		dI	12 22 43.14	14 18 08.9		21.53	-1.14	-0.71	20.48	0.78	2	-10.85
419**		dE	12 30 24.48	12 47 34.5		20.60	-1.44	-0.62	20.40	0.83	1	-10.84
420		dE/I	12 18 34.86	18 35 49.3		21.15	-1.29	-0.62	20.45	0.82	2	-10.84

Table 2 – *continued*

(1) ID	(2) Name	(3) Type	(4) α (J2000)	(5) δ (J2000)	(6) V_h (km s $^{-1}$)	(7) $B(6)$	(8) ICP	(9) OCP	(10) B_T	(11) P	(12) Class	(13) M_B
421		dI	12 34 36.14	11 04 23.1		21.20	-1.13	-0.56	20.47	0.82	2	-10.80
422		VLSB	12 33 52.45	12 07 01.6		21.85	-1.24	-0.84	20.50	0.99	1	-10.80
423		dE/I	12 49 23.81	11 12 06.6		21.25	-1.22	-0.54	20.50	0.79	2	-10.78
424		dI	12 43 21.77	11 00 18.8		21.55	-1.18	-0.72	20.52	0.75	2	-10.75
425		dE/I	12 32 25.52	12 08 54.7		21.45	-1.20	-0.65	20.56	0.83	2	-10.72
426		dI	12 32 33.52	12 47 22.8		21.55	-1.20	-0.51	20.57	0.82	2	-10.69
427		dE/I	12 38 06.47	12 17 51.8		21.44	-1.16	-0.50	20.66	0.85	2	-10.68
428**		VLSB	12 33 51.84	12 41 49.1		21.44	-1.54	-1.21	20.65	1.00	1	-10.63
429**		dE/I	12 42 45.07	11 22 43.1		20.29	-0.25	-0.31	20.64	0.00	2	-10.62
430		VLSB	12 25 55.25	18 20 16.9		21.91	-1.45	-0.80	20.69	0.98	1	-10.61
431		dE	12 35 20.57	11 06 43.2		21.28	-1.26	-0.44	20.75	0.85	2	-10.51
432		dE/I	12 38 00.41	11 34 29.6		21.60	-1.13	-0.60	20.80	0.83	2	-10.51
433		dE/I	12 20 33.16	16 43 54.5		21.66	-1.13	-0.64	20.79	0.82	2	-10.48
434		dI	12 35 38.18	12 40 53.5		21.57	-1.24	-0.52	20.92	0.84	2	-10.43
435		VLSB	12 32 37.82	11 24 44.1		22.41	-1.19	-0.87	20.98	1.00	1	-10.41
436		dE/I	12 29 09.32	12 29 43.6		21.50	-1.34	-0.48	20.90	0.86	2	-10.35
437		dI	12 46 46.51	11 39 19.5		21.39	-1.36	-0.41	20.91	0.89	1	-10.35
438		dI	12 33 11.99	11 12 53.6		21.59	-1.10	-0.45	21.04	0.85	2	-10.35
439		dI	12 19 34.95	17 14 24.4		21.51	-1.12	-0.44	20.96	0.86	2	-10.31
440		dE/I	12 30 01.86	12 56 52.8		21.40	-1.03	-0.39	20.93	0.88	2	-10.31
441		dI	12 29 01.10	12 33 32.8		21.69	-1.16	-0.54	21.01	0.84	2	-9.67
442		dI	12 38 26.36	11 39 11.2		21.98	-1.12	-0.62	21.15	0.91	2	-9.59
443		dE/I	12 33 24.74	12 24 11.3		21.66	-1.07	-0.42	21.16	0.90	2	-9.56
444		dI	12 38 47.03	12 14 17.0		21.97	-0.89	-0.54	21.28	1.00	2	-9.52
445		dE/I	12 35 29.48	12 40 59.6		21.84	-1.12	-0.47	21.27	0.86	2	-9.51
446**		dE/I	12 42 12.67	12 22 16.5		21.71	-1.04	-1.00	21.36	1.00	2	-9.38
447		dI	12 51 17.12	12 10 14.5		22.06	-1.33	-0.52	21.40	1.00	2	-9.33
448		dE	12 32 47.48	11 18 06.8		21.97	-1.25	-0.28	21.66	1.00	2	-9.17
449		dE/I	12 49 30.24	11 13 05.9		22.34	-0.98	-0.58	21.80	0.98	2	-8.93

(i) The Virgo images were taken under a variety of seeing conditions. In the worst conditions the seeing was 3 arcsec. Galaxies in these images are slightly bigger and have less concentrated light profiles than galaxies in Virgo images taken under more typical conditions or in the background fields. We simulated galaxies and found that differences in P caused by seeing variations of this magnitude were negligible relative to other uncertainties; this is not surprising given that we are selecting low surface brightness galaxies that are typically much bigger than 3 arcsec.

(ii) Occasionally galaxies of interest fell on cosmic rays or chip defects. Since all candidates on the images are being studied by eye (to assign ratings), it was easy to identify when this happened. In these cases, the cosmic ray or defect was removed by interpolation using the IRAF IMEDIT task. All subsequent analysis was performed on the repaired images.

(iii) Some objects appeared in more than one exposure (the fields in Fig. 1 overlap). Occasionally these had very different concentration parameters from each other: one or the other might have fallen near a chip edge or defect. In this event we only used the measured parameters for the undamaged image. Otherwise data from both exposures were combined and used.

(iv) Often, on the initial detection pass, very low surface brightness objects were detected as several separate objects, each centred on a local noise peak. These objects were identified upon inspection and all objects except the single object centred on the galaxy centre were removed from the catalogue. The subsequent photometry was performed on this single object.

(v) Similarly, on occasion where galaxies (typically dwarf irregular) had off-centre hotspots, the initial detection algorithm

identified objects centred on these hotspots. In these cases, we re-centred the galaxy and performed subsequent photometry on this re-centred object.

(vi) Many objects have close companions that were identified upon inspection. The $B(6)$ concentration parameters for such objects have little meaning. Total magnitudes were determined individually as described in Section 4.

(vii) A few objects, such as IC 3483 (heliocentric velocity 108 km s $^{-1}$), fell in the gaps between the CCDs and are missing from our catalogue, even though they are in the area outlined in Fig. 1. This may also have happened for some lower-luminosity galaxies. The survey area we quoted in the previous section includes a (small) correction for area lost in this way.

(viii) Two very high surface brightness galaxies that look like background galaxies but have velocities that place them in the Virgo Cluster were included (a search for all objects with known velocities was performed using the NASA/IPAC Extragalactic Data Base, hereafter NED). The two objects were VCC 1313 (a blue compact dwarf) and VCC 1627 (a compact low-luminosity elliptical galaxy). Both lack a detectable low surface brightness halo at large galactocentric radius, normally a defining characteristic of low-luminosity galaxies. Such galaxies are not common for $B < 17$ ($M_B < -14$), but were they to be very common at $B > 17$, they would be missing from our sample and the LF that we determine would be incomplete at the faint end. We do not, however, regard this possibility as likely and return to this issue in Section 7.

(ix) Low surface brightness features likely to be associated with luminous galaxies (whether cluster members or background) were judged not to be independent galaxies and were removed from the

catalogue. For example, there is a large stream of material flowing from the background galaxy pair IC 3481/A (heliocentric velocity $\sim 7000 \text{ km s}^{-1}$) that we excluded because it is likely to be associated with these galaxies and not the Virgo Cluster. Similarly, M100 in the cluster is surrounded by considerable amounts of low surface brightness material that we assume is associated with that galaxy and is not made up of individual cluster dwarfs.

(x) For low surface brightness objects that fell near the edge of a CCD, it was sometimes not clear whether the centre of the object fell on the CCD or not and it is possible that we included objects that we should have rejected under point (vii) above. However, this happened rarely enough that it is not a significant source of error.

4 GALAXY SAMPLE

Our sample consists of 449 galaxies of which we rated 102 as ‘0’, 220 as ‘1’ and 127 as ‘2’. Images of typical galaxies in our sample are presented in Fig. 3.

The galaxy sample is presented in Table 2. There we list the following.

(1) The galaxy identification number, ranked by M_B . A single star by the name means that we identified a companion close enough that it would affect photometry measurements at radii $6 < r < 12$ arcsec from the galaxy centre but not at $r < 6$ arcsec. Measurements of the OCP and P for these galaxies are consequently not an indicator of the galaxy properties. A double star by the name means that we identified a companion close enough that it would affect photometry measurements at $r < 6$ arcsec. Measurements of the $B(6)$, the ICP, OCP and P for these galaxies are consequently not an indicator of the galaxy properties for these objects.

(2) The galaxy name as appearing in the NASA/IPAC Extragalactic Data Base. The VCC numbers refer to the Virgo Cluster Catalog [Binggeli et al. (1985) and the IBM88 designations refer to objects studied by Impey et al. (1988)].

(3) The galaxy type, taken either from the Revised Third Catalog (RC3; de Vaucouleurs et al. 1991) of galaxies, the VCC, or our own inspection of the images. The designations are E = elliptical galaxy, S0 = lenticular galaxy, S = spiral galaxy, with the following letters indicating the Hubble type, BCD = blue compact dwarf, VLSB = very low surface brightness galaxy (these were labelled as such based on a visual inspection and tended to correspond to galaxies having average surface brightnesses within an aperture of radius 12 arcsec fainter than $26.5 B \text{ mag arcsec}^{-2}$; at the very faint end of the sample, this kind of classification becomes tenuous and the distinction between VLSB galaxies and other dwarfs is purely subjective), dE, dwarf elliptical/dwarf spheroidal galaxy; dS0, dwarf lenticular; dI, dwarf irregular. The notation dE/I means that we were unable to identify from the images what the type of dwarf galaxy was.

(4) Right ascension.

(5) Declination.

(6) Heliocentric radial velocity. See NED for the original references.

(7) $B(6)$, the apparent blue magnitude measured within a 6-arcsec radius.

(8) The ICP, or inner concentration parameter, defined as in Section 3. This is not meaningful for objects with companions.

(9) The OCP, or inner concentration parameter, defined as in Section 3.

(10) The total apparent blue magnitude B_T . For the very brightest

galaxies ($B_T < 15$) we adopt the RC3 values when available, since these galaxies normally extended beyond the edges of a single CCD chip in our data. For the majority of galaxies we computed B_T from the magnitudes within a 12-arcsec aperture, corrected for light at large radius that has fallen below the sky using the prescription of Tully et al. (1996). This method requires us to assume an azimuthally averaged exponential profile $B(r) = -2.5 \log_{10} \int_0^r I_0 \exp(-r/h) 2\pi r dr$, to derive the scalelength h and central intensity I_0 from a fit at radii below 12 arcsec, and then to compute $B_T = B(\infty)$. We estimate errors of ~ 0.2 mag in this analysis, derived from comparing B_T values derived from objects appearing in more than one image. More serious systematic errors might result if the light at large radii that fell below the sky is *not* exponential or follows a different exponential law from what we derive by fitting at smaller radii; however, there is no evidence for this phenomenon happening in the typical dwarf galaxies observed by Binggeli & Cameron (1991). The only galaxy where we did not use an exponential profile for this extrapolation was the elliptical galaxy NGC 4486A, where we used an $r^{1/4}$ -law extrapolation (de Vaucouleurs 1948). For objects labelled ‘**’ the companions prevent a fit from giving meaningful h and I_0 values (this effect was not significant for objects labelled ‘*’), so the B_T values were derived individually by identifying a symmetry axis and computing the flux from the part of the galaxy on the other side of this axis from the companion and multiplying this flux by 2. Additionally, VCC 2062 was treated this way, since the light from this irregular galaxy peaks so far away from its centre.

Since we will ultimately consider the LF for the entire magnitude range $-22 < M_B < -11$, it is important to consider how the RC3 B_T magnitudes used at very bright magnitudes compare with the B_T magnitudes we derive from the exponential fitting method at fainter magnitudes. We made the transition from using the different types of measurements at $B_T = 15$. For objects within 0.5 mag of this transition magnitude, the difference Δ between the B_T values obtained from the RC3 and those we derive is $\Delta = 0.16 \pm 0.10$ mag, which is smaller than our expected uncertainty in the B_T values we derive (see above) and much smaller than the binwidth (1 mag) we use to compute the LF. The

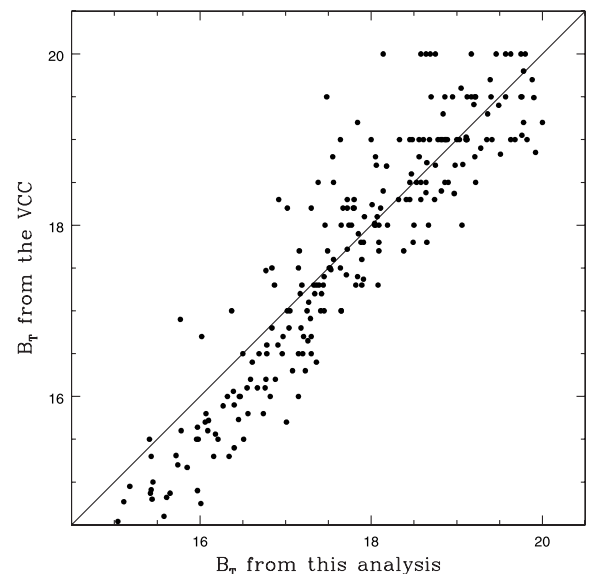


Figure 4. Comparison between the B_T magnitudes derived in the present analysis and those from the VCC (Binggeli et al. 1995).

RC3 mag will be more accurate for brighter galaxies, but even at the faintest limits that we use them, they therefore appear to be accurate enough for our purposes.

It is also instructive to compare our derived magnitudes with the magnitudes in the VCC for all dwarfs common to both catalogues (see Fig. 4). For $M_B = -16$ to -11 we have 318 galaxies, 232 (73 per cent) of which are VCC members. Some are not (such as UGC 7346) since they are not in the VCC fields – in this study we observed fields beyond the northern extent of the VCC (see Fig. 1) survey area. Some are not VCC members (such as the IBM88 objects) because their surface brightnesses were too low to be included in that sample. From Fig. 3, we see that the VCC magnitudes brighter than $B_T = 18$ tend to be systematically brighter than ours by approximately half a magnitude [a similar,

albeit smaller, effect was noted by Binggeli et al. (1985) who compared the VCC magnitudes with the B_T values derived by Ichikawa, Wakamatsu & Okamura 1986]. Fainter than $B_T = 17$ there is no systematic offset between the VCC magnitudes and ours [the mean difference is $B_T(\text{present}) - B_T(\text{VCC}) = -0.01 \pm 0.04$ mag].

(11) the nominal probability P of being a cluster member, as determined in Section 3 from a comparison with the background fields.

(12) Galaxy rating according to the scheme described in Table 1.

(13) Absolute magnitude $M_B = B_T - 31.15 - A_B$, where the Galactic extinction A_B values (typically 0.1–0.2 mag) are adopted from NED and based on the measurements of Schlegel et al. (1998). We adopt a distance modulus of 31.15 for the Virgo Cluster (Tonry et al. 2001).

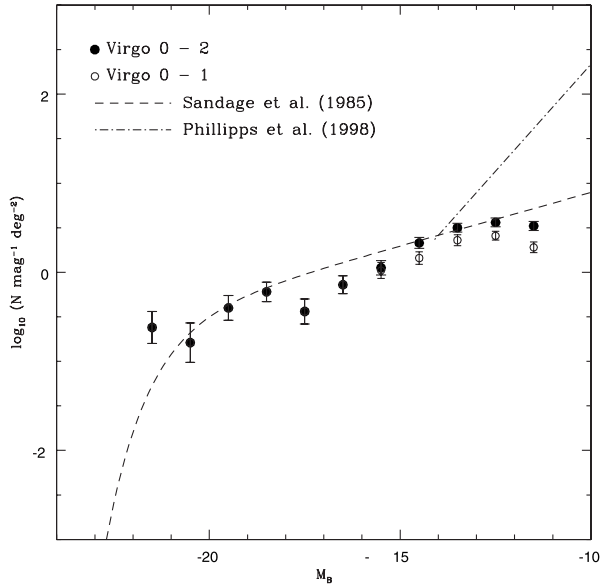


Figure 5. The luminosity function for the Virgo Cluster sample. Filled circles represent the luminosity function for all galaxies rated 0–2. Open circles represent the luminosity function for all galaxies rated 0–1. The dashed line represents the Schechter (1976) function fit of Sandage et al. (1985) for all galaxies with $B < 20$ ($M_B < -11.1$), including their incompleteness corrections. The dotted-dashed line represents the power-law fit of Phillipps et al. (1998a) to their inner-area sample for all galaxies with $15.5 < R < 20$ ($-15.6 < M_R < -11.1$), scaled horizontally assuming $B - R = 1.5$. The two fits are scaled vertically to have the same number of galaxies as our current sample at $M_B = -14$.

5 THE LUMINOSITY FUNCTION

The luminosity function for our galaxy sample is presented in Fig. 5 and Table 3.

Our luminosity function is similar to that derived by Sandage et al. (1985). This is not surprising given that most of our galaxies were in the Virgo Cluster Catalog (except at the faint end where our survey was deeper or the VCC is incomplete) and the fact that our total magnitudes for the VCC galaxies (Fig. 4) are similar to the VCC ones.

However, the difference with the results of Phillipps et al. (1998a) is very marked. Our survey reached a similar depth to theirs, but in the faintest three magnitude bins (Fig. 5) we found far fewer galaxies per bright galaxy. In our faintest bin we found only 83 galaxies, whereas given the inner area LF of Phillipps et al. (1998a) normalized to our data at $M_R = -14$ (the bright end of the magnitude range over which they fit the LF), we would predict ~ 1000 galaxies. Additionally, we found that our LF steepened at brighter magnitudes: $M_B = -16.5$ compared with $M_B = -13.5$ ($M_R = -15$) for the inner area LF of Phillipps et al. (1998a). This discrepancy can be explained if the sample of Phillipps et al. (1998a) suffers from severe background contamination at the faint end. Alternatively, it could mean that our sample is incomplete (it would need to be incomplete at approximately the 90 per cent level) at the faint end relative to the sample of Phillipps et al. (1998a). Such incompleteness could follow from us (and Sandage et al. 1985) erroneously rejecting the higher surface brightness cluster members from our sample because we think that they are background galaxies. We will return to this point in Section 7.

Table 3. The Virgo Cluster luminosity function.

M_B range	Number		$\log_{10}(N_{\text{gal}} \text{mag}^{-1} \text{deg}^{-2})$	α		
	0+1+2	0+1		0+1	0+1+2	0+1
$-22 < M_B < -21$	6	6	-0.62 ± 0.18	-0.62 ± 0.18		
$-21 < M_B < -20$	4	4	-0.79 ± 0.22	-0.79 ± 0.22	-1.32 ± 0.28	-1.32 ± 0.28
$-20 < M_B < -19$	10	10	-0.40 ± 0.14	-0.40 ± 0.14	-1.65 ± 0.28	-1.65 ± 0.28
$-19 < M_B < -18$	15	15	-0.22 ± 0.11	-0.22 ± 0.11	-0.95 ± 0.25	-0.95 ± 0.25
$-18 < M_B < -17$	9	9	-0.44 ± 0.14	-0.44 ± 0.14	-1.15 ± 0.19	-1.15 ± 0.19
$-17 < M_B < -16$	19	19	-0.12 ± 0.10	-0.12 ± 0.10	-1.64 ± 0.19	-1.60 ± 0.20
$-16 < M_B < -15$	31	29	0.10 ± 0.08	0.07 ± 0.08	-1.58 ± 0.14	-1.37 ± 0.15
$-15 < M_B < -14$	55	38	0.34 ± 0.06	0.18 ± 0.07	-1.49 ± 0.11	-1.37 ± 0.12
$-14 < M_B < -13$	79	57	0.50 ± 0.05	0.36 ± 0.06	-1.26 ± 0.09	-1.27 ± 0.11
$-13 < M_B < -12$	90	64	0.56 ± 0.05	0.41 ± 0.05	-1.02 ± 0.09	-0.90 ± 0.11
$-12 < M_B < -11$	83	47	0.52 ± 0.05	0.28 ± 0.06		

The values of α in this table come from power-law fits over a 3-mag range centred on the middle of the magnitude range for the appropriate entry.

Finally, we note that the final two points in the LF of Phillipps et al. (1998a) may have been overestimated by a factor of 2, meaning that a faint-end slope of $\alpha \sim -1.9$ may be appropriate for that data set, not $\alpha \sim -2.2$ (S. Phillipps, private communication). This would make the inconsistency with the current data set smaller.

Table 3 also gives α , the logarithmic slope of the luminosity function, at each absolute magnitude. The curvature in the LF is real and statistically significant: a single value of α , independent of absolute magnitude, is highly inconsistent with the data. Neither a power-law or a Schechter (1976) function provides a satisfactory fit to the data over any appreciable magnitude range. That we are able to make such a statement follows from the small (Poisson) error bars in Table 3, which in turn follows from the large number of galaxies in our sample. It is interesting, however, that the average value of α fainter than $M_B = -18$ is $\alpha = -1.35$, close to the faint-end slope that Sandage et al. (1985) found for the VCC. Over the magnitude range $-17 < M_B < -14$ the average slope is more like $\alpha = -1.7$, similar to the slope in the sample of Phillipps et al. (1998a) over this magnitude range, allowing for a $B - R$ colour of 1.5.

Over the half-magnitude interval $-11 < M_B < -10.5$, only have 29 galaxies classified 0 + 1 + 2, implying that the logarithm of the LF at $M_B = -10.75$ is $0.37 \pm 0.08 \text{ mag}^{-1} \text{ deg}^{-2}$. This could be a sign of a weak turnover in the LF, but there are other ways to explain the paucity of galaxies in this interval. For example, the sample could be incomplete at these faint levels (only a modest amount of incompleteness would be required to generate this kind of feature in the LF. This kind of incompleteness could follow from the reduced dynamic range in surface brightness over which we classify galaxies 1 or 2 at these faint levels (Trentham & Tully 2001). Galaxies with surface brightnesses a little brighter than $\sim 27 B \text{ mag arcsec}^{-2}$ will be missing from the sample because they are indistinguishable from field galaxies of the type seen in large number in our blank fields. Galaxies with surface brightnesses a little fainter than $\sim 27 B \text{ mag arcsec}^{-2}$ will also be missing from the sample because they are not detected above the sky. Consequently, at these very faint levels our sample may be incomplete.

6 PROPERTIES OF GALAXIES

The contribution to the total galaxy LF from galaxies of different morphological types is presented in Fig. 6.

At the faint end, the vast majority of galaxies in the sample are dE galaxies, as identified on morphological grounds. The structural parameters (see Fig. 7) and colours (see Fig. 8) of the galaxies are consistent with this interpretation. Many of these dE galaxies are nucleated, but few of the dIrr galaxies in the cluster are. This suggests either that the two kinds of galaxies form in different ways (despite the similarity in their scaling laws) or that dIs are dEs in formation and that the nucleus is the final part to form. In the faintest three bins approximately one-third of the galaxies are rated 2. These tended to be high surface brightness dEs that could conceivably be background late-type galaxies. Even excluding these galaxies, dEs are still the dominant types at the faintest magnitudes.

The VLSB galaxies never contribute significantly to the total luminosity function. This result cannot be compared directly with the measurements of VLSB galaxies in Fornax of Kambas et al. (2000) because we require a somewhat lower surface brightness (see Section 4) than do Kambas et al. (see Section 2.2 of their paper) to call a galaxy a ‘VLSB’. Nevertheless, we do find far

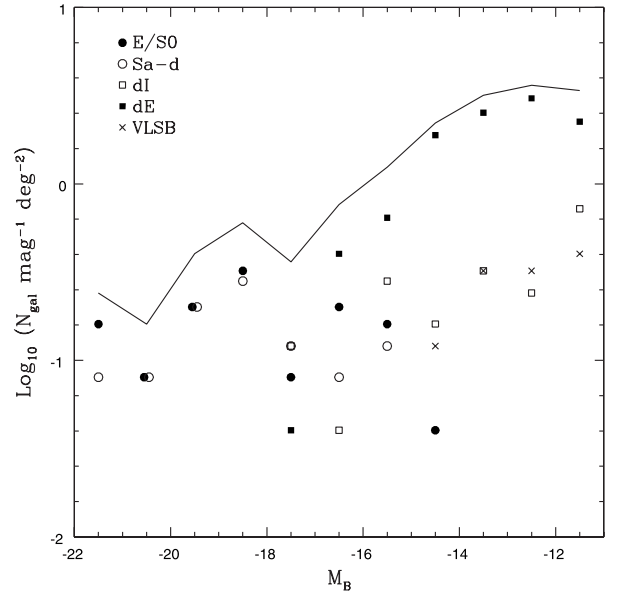


Figure 6. The luminosity function segregated by morphological type (galaxies classified dE/I or dS0 in Table 2 are grouped with the dEs here and galaxies classified as BCD in Table 2 are grouped with the dIs here). The line indicates the total luminosity function, as presented with uncertainties in Fig. 5.

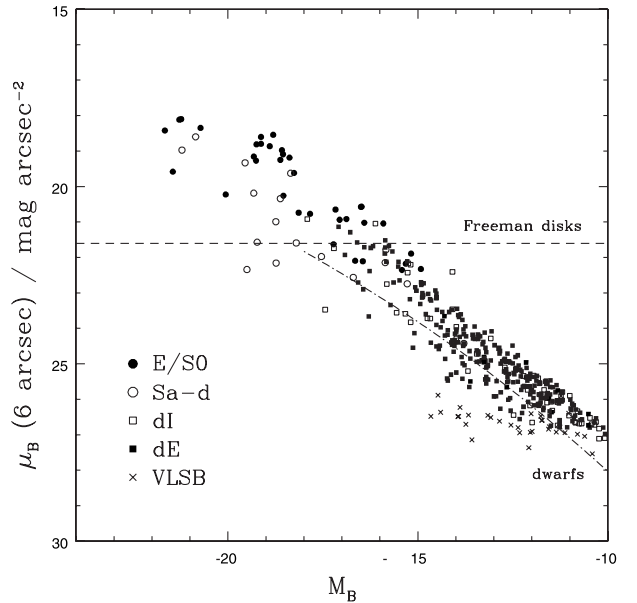


Figure 7. Surface brightness μ_B versus absolute magnitude M_B for the sample galaxies, segregated by morphological type (galaxies classified dE/I or dS0 in Table 2 are indicated as dE here and galaxies classified as BCD in Table 2 are indicated as dI here). Surface brightnesses are the average values measured within circular apertures of radius 6 arcsec. Galaxies with companion stars or other galaxies projected within this aperture are not included. The value of $\mu_B = 21.6 \text{ arcsec}^{-2}$ for galaxy discs (Freeman 1970) is represented by the dashed lines; many spiral galaxies are above this line owing to the presence of a bulge. The dotted-dashed line represents typical values (e.g. Binggeli & Cameron 1991; Binggeli 1994) for exponential dwarf galaxies; many dE galaxies are above this line owing to the presence of a nucleus and many dI galaxies are above this line owing to the presence of star-forming knots within the aperture. Both dwarf ellipticals and dwarf irregulars have similar scaling laws (Binggeli & Cameron 1991).

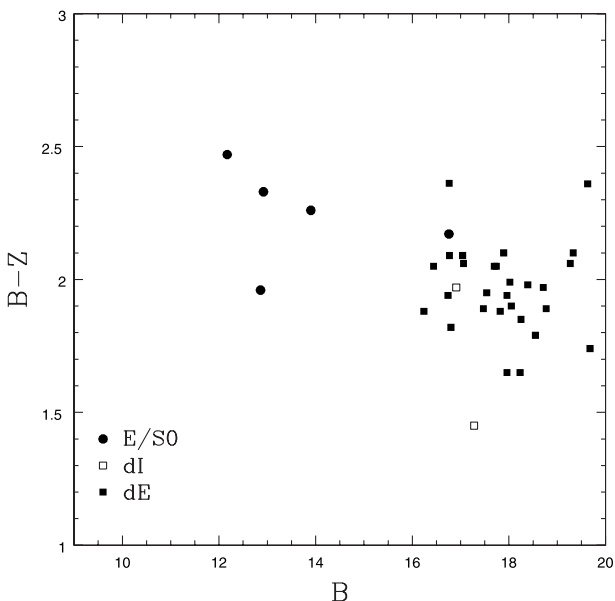


Figure 8. Colour–magnitude diagram for the 35 galaxies that were in the nine fields observed in Z and bright enough to be detected at the 5σ level in the Z data. All magnitudes were measured in an aperture of radius 6 arcsec, which was large enough that the effect on the colours caused by differential seeing between the B and Z data was negligible. Only galaxies with total absolute magnitudes of $M_B < -12$ are shown. Fainter (and consequently lower surface brightness galaxies) were not detected above sky in the Z-band data, particularly since that data suffers from serious fringing, typically 6 per cent of the sky (<http://www.ast.cam.ac.uk/~wfsur/defringing.html>). Dwarf elliptical galaxies have $B - Z$ colours of 1.7–2.2 given the spectral energy distributions of Trentham et al. (1998b).

fewer low surface brightness galaxies that are likely to be cluster members than they did in Fornax. The Virgo Cluster has crossing times shorter than one-tenth of a Hubble time (see table 1 in Tully et al. 1996) so galaxies here must have undergone many galaxy–galaxy interactions. It is therefore surprising that diffuse VLSB galaxies can survive. Perhaps this is evidence that small galaxies in clusters have substantial dark matter haloes.

We also notice a paucity of galaxies having exactly $M_B = -17.5$. A dip in the LF at this absolute magnitude was also seen in the Coma Cluster (Trentham 1998a) and in the NGC 1407 Group (Trentham & Tully 2001), another dense knot of early-type galaxies. This is also the magnitude where the cluster population changes from being dominated by high surface brightness giant elliptical galaxies to low surface brightness dwarf elliptical galaxies (see Fig. 6). These two types of stellar system have very different structural parameters (see fig. 1 of Binggeli 1994), implying that they had very different formation mechanisms.

7 CAVEATS: INCOMPLETENESS AND CONTAMINATION

We expect to be missing two sorts of galaxies from our sample: galaxies with very low surface brightnesses that are never visible above the sky and cluster galaxies with high surface brightness that we reject from the sample because we think that they are background galaxies. We will argue that neither of these are likely to be a serious problem with the current data set.

Galaxies with very low surface brightnesses ($\mu_B > 27$ mag arcsec $^{-2}$ within an aperture of radius 6 arcsec) are unlikely to be very common for $M_B < -11$. Evidence supporting this is as follows.

(i) Such galaxies were not found in deep images of the Virgo Cluster core (Trentham & Tully 2001) taken with the 8-m Subaru Telescope. Those images reached very deep surface brightness limits (~ 28 R mag arcsec $^{-2}$, equivalently ~ 29 B mag arcsec $^{-2}$ for dwarf galaxies) but did not uncover any additional galaxies with $M_B < -11$ that were missing from the current survey because their surface brightnesses were too low.

(ii) In Fig. 7, the points at the faint M_B end do not cluster right at the very limit of detection (27 mag arcsec $^{-2}$). Most have higher surface brightnesses (although a few have $\mu_B \sim 27$ mag arcsec $^{-2}$). This in turn suggests that there do not exist large numbers of galaxies with surface brightnesses just below this limit.

(iii) No local galaxies with $M_B < -11$ are known with such low surface brightnesses, although a small number of bulge-dominated galaxies have disc components this faint or fainter, e.g. Malin 1 (Bothun et al. 1987) and GP 1444 (Davies, Phillipps & Disney 1988). The giant galaxy population in the Virgo Cluster is similar to that in the local Universe in terms of galaxy structural parameters, so it is reasonable to expect that the dwarf galaxy population is similar too and that this absence in the field of dwarfs with $M_B < -11$ and $\mu_B > 27$ mag arcsec $^{-2}$ extends to the Virgo Cluster.

Dwarf galaxies with very high surface brightnesses lacking an extended diffuse light component are another potential source of contamination because they would be rejected from the current sample since they look like luminous background galaxies (they have low P values). These can either be blue H II galaxies such as Markarian 1460 (Trentham, Tully & Verheijen 2001b) or red compact dwarfs such as M32 in the Local Group or UGC 6805 in the Ursa Major Cluster (Tully & Verheijen 1997). In the Virgo Cluster, VCC 1313 is an example of the former and VCC 1627 of the latter; were the velocities of these two objects not known we would have classified them as background objects. For $M_B < -14$ only these two galaxies of this type were identified. Were such objects to exist with $M_B > -14$, our sample could be incomplete. The fraction of detected galaxies that we assign to the cluster on surface brightness grounds ranged from close to 1 at the bright end of the sample ($B < 15$) to ~ 1 per cent at the faint end ($B \sim 20$).

This is an important change of emphasis from what has been thought classically of as the main uncertainty in studies of cluster luminosity functions. Previously, the results from this kind of study were open to question because many very low surface brightness galaxies could be missing from the sample since they are never detected above the sky. This is no longer a worry because deep surveys such as this one and that of Trentham & Tully (2001) are not uncovering large numbers of LSB galaxies that were missing in shallow ones. Instead, the major concern is now that the sample may be missing many *high* surface brightness galaxies that we have culled from the sample because we think that they are background galaxies.

We do not, however, regard this as a serious worry. First, at the bright end of our sample, where velocity measurements are available, galaxies with high surface brightnesses are rare (VCC 1313 and VCC 1627 are the exceptions). Secondly, high surface brightness galaxies do not appear to be present in substantial number in the Fornax Cluster else they would have been seen in the spectroscopic survey described by Drinkwater et al. (2000a; this work uses the 2dF spectrograph on the Anglo-Australian Telescope). Small numbers of compact galaxies with early-type spectra were discovered in that survey (Drinkwater et al. 2000b, Phillipps et al. 2001), but they are too rare to contribute

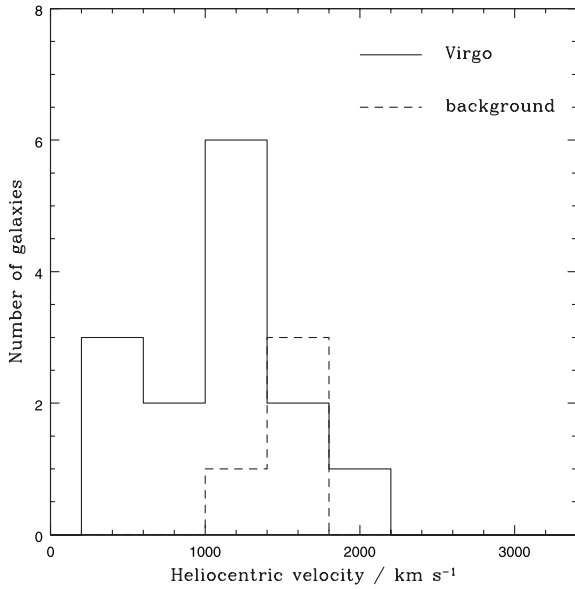


Figure 9. Histogram of all galaxies in the Virgo and North Galactic Cap background regions listed in the *Nearby Galaxies Catalog* (Tully 1987) and having $M_B < -16$. The peak in the Virgo sample at $\sim 1000 \text{ km s}^{-1}$ represents the Virgo Cluster.

significantly to the Fornax LF. Given the similarities between the Virgo and Fornax Clusters, we do not expect this to be a major source of incompleteness in the current study.

Another potential problem is that the Virgo data, but not the North Galactic Cap background data, are contaminated by an anomalously large number of nearby galaxies at comparable distances to the Virgo Cluster that are not bound to the cluster (the P values computed by comparing the numbers of low surface brightness galaxies in the two data sets would then be too high). The contaminating galaxies would need to be nearby or else they would not look like Virgo Cluster galaxies. We regard such a possibility as unlikely owing to the paucity of luminous galaxies with velocities between ~ 2000 and 3000 km s^{-1} in the Virgo data set (see Fig. 9). Were such objects numerous, the satellite populations of these luminous galaxies (only very luminous galaxies are listed in the *Nearby Galaxies Catalog* at these distances and so would be included in Fig. 9) could look like Virgo Cluster members (satellites of galaxies at higher velocities would be too small).

8 THE TOTAL OPTICAL LUMINOSITY IN GALAXIES

In Fig. 10, we present the contribution from galaxies of different types and absolute magnitudes to the total optical luminosity of the region of the Virgo Cluster that we surveyed. The total luminosity in our sample is $6.3 \times 10^{11} L_{\odot B}$, corresponding to a luminosity surface density of $5.6 \times 10^{10} L_{\odot B} \text{ Mpc}^{-2}$. For comparison, for the VCC, Sandage et al. (1985, converted to the distance scale used elsewhere in this paper) measure a luminosity surface density of $3.4 \times 10^{10} L_{\odot B} \text{ Mpc}^{-2}$ averaged over the central six degrees of the cluster. That our number is slightly higher follows from the fact that by proportion our survey covers more high-density areas within the cluster than does the VCC (see Fig. 1).

Fig. 10 shows that only a small proportion (less than one-tenth) of the total optical luminosity of the cluster is in dwarf galaxies.

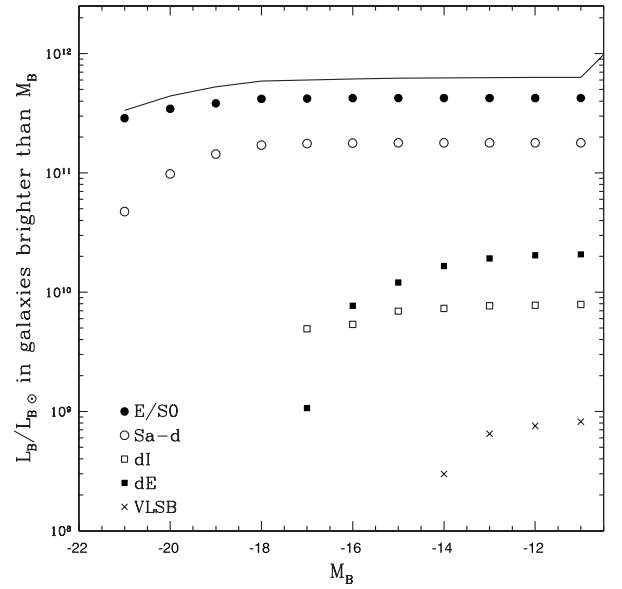


Figure 10. The total blue luminosity in galaxies brighter than M_B , segregated by morphological type (galaxies classified dE/I or dS0 in Table 2 are indicated as dE here and galaxies classified as BCD in Table 2 are indicated as dI here). The line indicates the total luminosity in galaxies of all morphological types.

Even less is in the systems that we called VLSB galaxies. Were these galaxies not to be resolved as individual objects (as would be the case if they were in a more distant cluster), they would be observable only though their contribution to the diffuse intracluster light. This in turn implies that if the cluster LF does not vary strongly between clusters (see the next section), then the intracluster light in distant clusters (typically > 10 per cent of the total cluster light e.g. Melnick, Hoessel & White 1977; Thuan & Kormendy 1977; Scheick & Kuhn 1994; Vílchez-Gómez, Pelló & Sanahuja 1994) cannot be produced by the integrated light from low surface brightness dwarf galaxies and is more likely to be made up of stars tidally released from luminous galaxies within the cluster, perhaps via galaxy harassment (Moore et al. 1996).

9 COMPARISON WITH OTHER LFS

In Fig. 11 we compare the Virgo Cluster LF with the LFs for the Coma Cluster ($z = 0.023$), the Ursa Major Cluster, and the Local Group. In Table 4 we present the dwarf-to-giant ratios (DGRs) for the various environments. The DGR is a convenient way to parametrize the LF by a single number; note that different authors, e.g. Phillipps et al. (1998b), define this quantity in a different way. The Virgo LF presented here and the Ursa Major LF (Trentham, Tully & Verheijen 2001a) are somewhat more tightly constrained than the Coma LF or the Local Group LF. For the Coma Cluster, the error bars are large owing to the need to determine the LF using a background subtraction and the large field-to-field variance of the background (the Coma Cluster is sufficiently distant that the faint dwarfs become smaller than the seeing and we cannot establish membership on the basis of morphology, as we did in the current study). For the Local Group, the error bars are large owing to Poisson counting statistics, since there are not many galaxies (there are only six galaxies – M31, the Milky Way, M33, IC 10, LMC and SMC – brighter than $M_B = -16$, for example).

There appear to be two types of galaxy luminosity function – one for evolved regions (where the elliptical galaxy fraction is

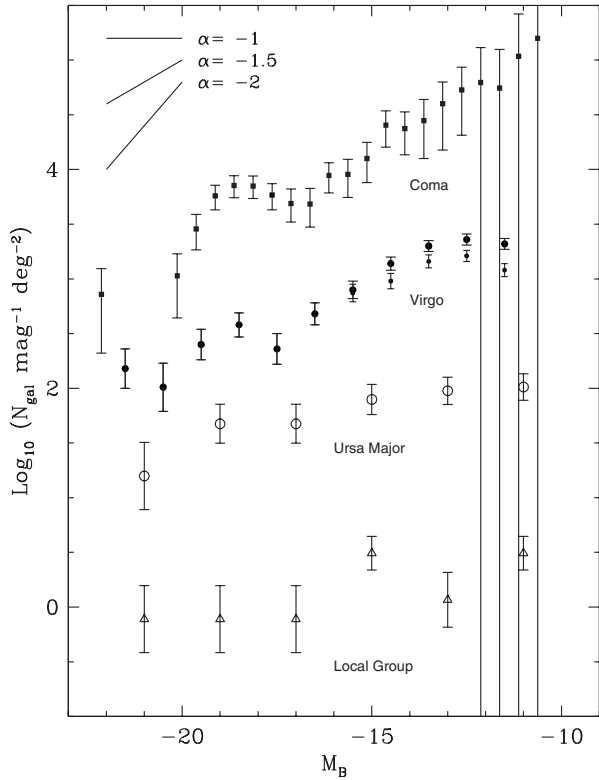


Figure 11. The B -band luminosity functions of the Local Group (open triangles), the Ursa Major Cluster (open circles), the Virgo Cluster (filled circles) and the Coma Cluster (filled squares). The Virgo Cluster data are from this work, shifted upward in the ordinate axis by 2.8 units (to permit clearer presentation). The Ursa Major data are from Trentham et al. (2001a). The Coma Cluster data are from Trentham (1998a), shifted upward in the ordinate axis by 1.5 units. The Local Group data are from the compilation of Irwin (http://www.ast.cam.ac.uk/~mike/local_members.html), adjusted to the B -band (where independent photometric measurements do not exist in the NED data base) assuming $B - V = 0.6$ for the Milky Way (see the template Sbc galaxy spectral energy distribution of Coleman, Wu & Weedman 1980), $B - V = 0.3$ for dwarf irregular galaxies (Coleman et al. 1980) and $B - V = 0.8$ for dE/dSph galaxies (Caldwell 1983). This Local Group data are shifted upward on the ordinate axis by 5.5 units.

high, the galaxy density is high and the crossing time is short) such as the Virgo Cluster and Coma Cluster and one for unevolved regions such as the Ursa Major Cluster and the Local Group. The Virgo and Ursa Major LFs represent natural prototypes for the two kinds of LF. The two LFs are inconsistent with each other at a high level of significance: the probability that the two LFs are drawn from a single distribution is $\ll 1$ per cent (reduced $\chi^2 = 22.2$ for five degrees of freedom). The reason that such a strong statement can be made follows from the small error bars for these two LFs in Fig. 11.

The evolved region LF appears to be characteristic of many clusters (see Trentham 1998c), even though for each cluster individually (such as Coma) the LF is poorly determined owing to the large uncertainties following a background subtraction. In the study of Trentham (1998c), the composite cluster LF was determined primarily from the LFs of Virgo (Sandage et al. 1985) and Fornax (Ferguson & Sandage 1988). The current LF, at least in all but the faintest one or two magnitude bins, can to some extent be seen as a verification of the Virgo LF of Sandage et al. (1985) and therefore the LF presented in Table 3 may be regarded

Table 4. Dwarf-to-giant ratios for various environments.

Environment	$N(-16 < M_B < -11)/N(M_B < -16)$
Coma	11.73 ± 6.76
Virgo (0 + 1 + 2)	5.36 ± 0.74
Virgo (0 + 1)	3.73 ± 0.53
Ursa Major	2.07 ± 0.67
Local Group	2.83 ± 1.35

Table 5. Faint-end slopes of spectroscopic field surveys.

Survey	α	Reference
Stromlo-APM	-1.0	Loveday et al. (1992)
Hawaii-Caltech	~ -1.25	Cowie et al. (1996)
Autofib	-1.1	Ellis et al. (1996)
LCRS	-0.7	Lin et al. (1996)
Sloan	-1.2	Blanton et al. (2001)
2dF + 2 MASS (near-IR)	-1.0	Cole et al. (2001)

as being valid (when scaled appropriately) for the majority of galaxy clusters. The main features of this evolved LF are the steep rise at $M_B = -16$ and the flattening faintward of $M_B = -14$. In the very centres of the richest clusters, the galaxy density is high enough that many dwarfs may be destroyed via cluster-related processes (e.g. Phillipps et al. 1998b; Adami et al. 2000; Boyce et al. 2001) and the rise at $M_B = -16$ is no longer observable there, but the Virgo Cluster is unlikely to be rich enough for this phenomenon to be important.

The unevolved (i.e. Ursa Major) LF is different from the evolved one in that it lacks the rise ($\alpha = -1.6$) at $M_B = -16$ and consequently generates a lower DGR. This LF ($\alpha \sim -1.1$ everywhere fainter than $M_B = -18$) is also appropriate for the Local Group (van den Bergh 1992, 2000) and for the true field, where the LF is determined spectroscopically [see Table 5; the values of α quoted by different authors are derived in different ways, for example, by Schechter (1976) function fits with different L^* , but none are as steep as the Virgo LF at around $M_B = -16$].

10 COMPARISON WITH THEORY

The discussion in the previous section can be summarized by the following results.

(1) The Virgo Cluster LF has a rise with $\alpha = -1.6$ at $M_B = -16$ that is not seen in the LFs of the field or unevolved environments such as the Ursa Major Cluster.

(2) This rise does not continue indefinitely towards fainter magnitudes and the Virgo Cluster LF is flat ($\alpha \sim -1.0$) between $M_B = -14$ and the limit of our survey at $M_B = -11$.

Furthermore,

(3) Dwarf galaxies are very deficient in the Virgo Cluster compared with the predictions of cold dark matter (CDM) theory (which predicts $\alpha \sim -2$ if light traces mass). The implication of (1) is then that dwarfs are even more deficient in the field.

Attempting to explain these three results in the context of galaxy formation models leads us to consider the following four physical processes.

(i) Tully (2001; see also Somerville 2001; Tully et al. 2002) has suggested that the dark haloes in dense, evolved environments

Table 6. Confrontation of observation with theory.

Process	Steepening of Virgo but not UMa/field LF at $M_B = -16$	Flattening of Virgo LF faintward of $M_B = -14$	Deficiency of dwarf galaxies compared with CDM predictions
Squelching	✓		✓
Supernova feedback		✓	✓
Cluster-related processes	✓		
Reducing $P(k)$ on small scales			✓

formed early in the history of the Universe, prior to reionization, but that the dark haloes in diffuse, unevolved environments assembled much later. The dark haloes in the evolved environments, such as the Virgo Cluster, could then collect gas that could later be turned into stars, but dark haloes in unevolved environments could not – in the terminology of Tully (2001), the formation of stars within these haloes would be ‘squelched’. The end result would be more dwarf galaxies per unit total mass in evolved environments than in unevolved environments, in agreement with result (1) above. In unevolved environments there would be very many dark matter haloes with no stars at all. Such a phenomenon is also seen in simulations (Chiu, Gnedin & Ostriker 2001).

(ii) Local feedback can generate a very low star formation efficiency in low-mass galaxies. Winds from a modest number of supernovae can expel a large fraction of the gas in small galaxies (Dekel & Silk 1986; Efstathiou 2000) because the galaxies have small potential wells. The consequence of this is that the final luminosity L of a small galaxy is very strongly decreasing function of the galaxy mass M so that M/L is a decreasing function of M . The LF that we measure is then much shallower than the mass function, which is what CDM predicts. That this process is important is suggested by the observation that the lowest-mass dwarf galaxies are heavily dark matter dominated (Wilkinson et al. 2002). This mechanism solves the discrepancy with CDM theory in a different way from squelching – it causes all low-mass galaxies to have very low luminosities, rather than turning off star formation altogether in all but very few low-mass haloes, leaving most low-mass haloes completely dark.

(iii) Cluster-related processes such as galaxy harassment (Moore et al. 1996) and tidal interactions during the early stages of cluster formation can form dwarfs (Barnes & Hernquist 1992) and may in part be responsible for the higher DGR in clusters. On the other hand, dwarfs formed in this way are not be expected to have dark haloes (and by implication, low surface brightnesses) such as the Virgo dwarfs we studied here; the dwarfs found in the simulations of Barnes & Hernquist do not have an appreciable dark matter content. Additionally, field and cluster dwarfs seem to have similar scaling laws (see Fig. 7), which would seem to argue that field and cluster dwarfs formed in a similar way to each other so that cluster-related processes are probably not the main mechanisms responsible for the difference in the evolved and unevolved LFs.

(iv) The number of dwarfs predicted by theory is very much lowered if the power spectrum $P(k)$ of primordial fluctuations is reduced on small scales. A general phenomenon of CDM theory is its success on large scales but its failure to reproduce observations on small (< 10 -kpc) scales – the two most serious failures are the dwarf galaxy deficiency studied here and the observed flatness of dark matter profiles in the centres of galaxies (e.g. Binney, Gerhard & Silk 2001). Possible mechanisms to remove small-scale power from the CDM fluctuation spectrum include making the dark

matter warm (Bode, Ostriker & Turok 2001) or self-interacting (Spergel & Steinhardt 2000, however, see Moore et al. 2000; Miralda-Escudé 2002). Studying the effects of these particular modifications to the dark matter power spectrum is currently an active field of study.

In Table 6 we summarize how the different physical processes can explain the observational results listed above. A check is placed in a column whenever an observational result follows from the inclusion of the physical process in question in models of galaxy formation.

11 FUTURE WORK

A natural extension of the current study is the determination of a deeper and more accurate Virgo Cluster LF. For $M_B < -11$, the main source of uncertainty in the current work is in establishing membership, not in counting statistics: the difference between the 0–1 and 0–2 points in Fig. 5 is greater than the error bars. Galaxies classified ‘2’, those for which we are uncertain about membership, tend to be the higher surface brightness ones in our sample, so it should be possible to establish distances to these either by measuring spectroscopic redshifts or from surface brightness fluctuations (Jerjen, Freeman & Binggeli 1998, 2000). For $M_B > -11$, deeper observations than those presented here over similarly large angular areas will be required. Such observations are now possible with the advent of mosaic CCDs on 8-m telescopes such as Suprime-cam on the Subaru Telescope (for deep observations of 1.2 deg^2 in the core of the Virgo Cluster, see Trentham & Tully 2001).

A spectroscopic survey of the Virgo Cluster could be extended to include significant numbers of compact, high surface brightness galaxies that we think are background, in the style of the Drinkwater et al. (2000a) Fornax survey. If the vast majority of such objects are indeed background objects, this would alleviate the concern highlighted in Section 7 that the current sample is heavily incomplete because of us rejecting such galaxies. Such a project is now feasible with the advent of wide-field multi-object spectrographs on large telescopes.

More detailed studies of the dwarf galaxies found in Virgo will also be of value in assessing the importance of the various physical processes at work during galaxy and cluster formations. The following observations should be of particular use.

(i) Measurement of colours and elemental abundances. These will constrain the star formation histories of the Virgo dwarfs, which in turn will allow a lower limit to be placed on the redshift at which gas was collected by small dark matter haloes. This is of importance in the context of the squelching picture described above. This squelching picture also relies on the very existence of these dark matter haloes around the Virgo dwarfs. In the long term,

studies of the kinematics of stars in the dwarfs will be required to verify this assertion.

(ii) H I measurements of the cold gas content of the Virgo dwarfs. Galaxies that have stayed any appreciable time in the Virgo Cluster would lose their gas via ram-pressure stripping from the cluster X-ray halo. Therefore, a large number of dwarfs in Virgo with H I would suggest that many Virgo dwarfs only recently entered the cluster. This would, in turn, argue against any physical process that requires Virgo dwarfs to have formed in the cluster or at very early times in small groups that quickly merged to form the cluster.

(iii) The location of dwarfs in the cluster. Do giants maintain their dwarf populations once they are in the Virgo Cluster, or do the dwarfs adopt orbits determined by the cluster potential? If the former were true, dwarfs would tend to cluster around giants. If the latter were true, they would be smoothly distributed throughout the cluster, with a radial density profile similar to that of the giant galaxies. The answer to this question will provide constraints on the dark matter structure of the Virgo Cluster.

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